WELCOME !

Astrophysics IV :

Stellar & Galactic Dynamics

Spring 2025

Dr. Yves Revaz
Laboratoire d’astrophysique
Observatoire de Sauverny
CH – 1290 Versoix

EPFL

Mailing List

- Use Moodle : moodle.epfl.ch

Anyone missing ?

About me

- MER at the Laboratory of Astrophysics
- Native from le Valais
- Former EPFL student
- Thesis in galactic dynamics (Prof. Pfenniger)
- Postdoc in Geneva, Paris and EPFL
- Faculty at EPFL since 2010

Research

- Formation and evolution of galaxies
- Galactic dynamics, galaxy clusters, dwarf galaxies
- Development of numerical tools (Gear, pNbody, Swift)
- Core Team Member of the Arrakihs Space mission
- Work package leader of WP2 of the Arrakihs Space mission
- Virtual reality
  - VIRUP: The Virtual Reality Universe Project
  - https://go.epfl.ch/virup

Contacts

- Email : yves.revaz@epfl.ch
- Cubotron (BSP) 323
- Observatoire de Sauverny, 351


Suisse
France
Berolle
Apples
Crissier
Bois-d'Amont
Bière
Prilly
Ecublens
Lausanne
Parc Jura vaudois
Gimel
Morges
Saint-Sulpice
Pully
Le Noirmont
Longirod
Marchissy
Aubonne
Etoy
Saint-Prex
Allaman
Rolle
Gilly
Arzier-Le Muids
Saint-Cergue
Genolier
Dully
Lac Léman
Gland
Gingins
Grens
Prangins
Nyon
Évian-les-Bains
Lugrin
Publier
Thollon-les-Mé
Saint-Paul-en-Chablais
Champanges
Thonon-les-Bains
Bernex
Yvoire
Excenevex
Anthy-sur-Léman
Divonne-les-Bains
Margencel
Sciez
Chens-sur-Léman
Coppet
VAUD
GENÈVE
Douvaine
Orcier
Perrignier
La Forclaz
Vacheresse
Loisin
Lullin
Bonnevaux
La Baume
Veigy-Foncenex
Brenthonne
Bons-en-Chablais
Corsier
Le Biot
Ferney-Voltaire
Collonge-Bellerive
Machilly
Bellevaux
Habère-Poche
Seytroux
Pregny-Chambésy
Saint-Cergues
Habère-Lullin
Saint-Jean-d'Aulps
Le Grand-Saconnex
Villard
Boëge
Genève
Ville-la-Grand
Annemasse
Bogève
Montriond
D1005
D902
D1206
D903
E62
1
5
9


Suisse
France
Bois-d'Amont
Berolle
Bière
Apples
Crissier
Prilly
Ecublens
Lausanne
Pully
Parc Jura vaudois
Gimel
Morges
Saint-Sulpice
Le Noirmont
Longirod
Marchissy
Aubonne
Etoy
Saint-Prex
Bourg
Allaman
Rolle
Arzier-Le Muids
Gilly
Saint-Cergue
Genolier
Dully
Gland
La Dôle
Lac Léman
Suisse
France
Suisse
Divonne-l
D1005
Ferney-Voltaire
Pregny-
Le Grand-Sac
Ge
le Cubotron


taine
Suisse
France
Berolle
Apples
Crissier
Bois-d'Amont
Bière
Prilly
Ecublens
Lausanne
Parc Jura vaudois
Gimel
Saint-Sulpice
Morges
Pully
Le Noirmont
Longirod
Marchissy
Aubonne
Etoy
Saint-Prex
ses
Bourg
Allaman
Arzier-Le Muids
Rolle
Gilly
Saint-Cergue
Suisse
France
Lac Léman
Genolier
Dully
a Dôle
Gland
Suisse
France
Gingins
Mei
Évian-les-Bains
Lugrin
Grens
Prangins
Publier
Thollon-les-Mé
Nyon
Saint-Paul-en-Chablais
Thonon-les-Bains
Champanges
Yvoire
Bernex
Divonne-les-Bains
Excenevex
Anthy-sur-Léman
Suisse
France
Margencel
Sciez
Chens-sur-Léman
La Forclaz
Vacheresse
ssy
Coppet
Orcier
VAUD
GENÈVE
Douvaine
Perrignier
Bonnevaux
Loisin
Lullin
La Baume
Abo
Veigy-Foncenex
Brenthonne
Ferney-Voltaire
Corsier
Bons-en-Chablais
Le Biot
Bellevaux
Collonge-Bellerive
Machilly
Habère-Poche
Seytroux
Pregny-Chambésy
Saint-Cergues
Le Grand-Saconnex
Habère-Lullin
Saint-Jean-d'Aulps
Villard
Boëge
Genève
Ville-la-Grand
Montriond
Annemasse
Bogève
D1005
D902
D1206
D903
E62
1
5
9


Suisse
France
Bois-d'Amont
Berolle
Apples
Crissier
Prilly
Ecublens
Lausanne
Pully
Bière
Parc Jura vaudois
Gimel
Morges
Saint-Sulpice
Le Noirmont
Longirod
Marchissy
Aubonne
Etoy
Saint-Prex
Allaman
Rolle
Saint-Cergue
Gingins
Grens
Divonne-les-Bains
VAUD
GENÈVE
Ferney-Voltaire
Pregny-Chambésy
Le Grand-Saconnex
Genève
Collonge-Bellerive
Corsier
Veigy-Foncenex
Loisin
Brenthonne
Bons-en-Chablais
Machilly
Saint-Cergues
Ville-la-Grand
Annemasse
Lullin
Bellevaux
Habère-Poche
Habère-Lullin
Villard
Boëge
Bogève
La Baume
Le Biot
Seytroux
Saint-Jean-d'Aulps
Bonnevaux
D1005
D1206
D903
D902
E62
1
5
9

Astrophysics @ EPFL Teaching

- **Astro I:** Introduction à l'astrophysique (Bachelor)
  - Jean-Paul Kneib
- **Astro II:** Bases physiques de l'astrophysique (Master)
  - Pascale Jablonka
- **Astro III:** Galaxy Formation and Evolution (Master)
  - Michaela Hirschmann
- **Astro IV:** Stellar and Galactic Dynamics (Master)
  - Yves Revaz
- **Astro V:** Observational Cosmology (Master)
  - Jean-Paul Kneib
- The Variable Universe (EDPY)
  - Richard Anderson
- **MOOC:**
  - The radio-sky I : Science and Observations
    Frédéric Courbin, Jean-Paul Kneib
  - Introduction à l'astrophysique
    Frédéric Courbin

Astrophysics @ EPFL Teaching

- **Astro I:** Introduction à l'astrophysique (Bachelor)
  - Jean-Paul Kneib
- **Astro II:** Bases physiques de l'astrophysique (Master)
  - Pascale Jablonka
- **Astro III:** Galaxy Formation and Evolution
  - Michaela Hirschmann
- **Astro IV:** Stellar and Galactic Dynamics (Master)
  - Yves Revaz
- **Astro V:** Observational Cosmology (Master)
  - Jean-Paul Kneib
- The Variable Universe (EDPY)
  - Richard Anderson
- **MOOC:**
  - The radio-sky I : Science and Observations
    Frédéric Courbin, Jean-Paul Kneib
  - Introduction à l'astrophysique
    Frédéric Courbin

In addition:

- TP4a
- TP4b
- Specialisation semester
- Master's project

Astrophysics @ EPFL
Research

Research group leaders :

- Jean-Paul Kneib
- Michaela Hirschman
- Pascale Jabonka
- Yves Revaz
- David Harvey
- Richard Anderson
- Emma Tolley

Research fields:

Galaxy Formation & Evolution

Cosmological parameters

Epoch of reionization

Dark energy

Dark matter

Astrophysics @ EPFL
Research

Research group leaders :	Jean-Paul Kneib Michaela Hirschman Pascale Jabonka Yves Revaz David Harvey Richard Anderson Emma Tolley

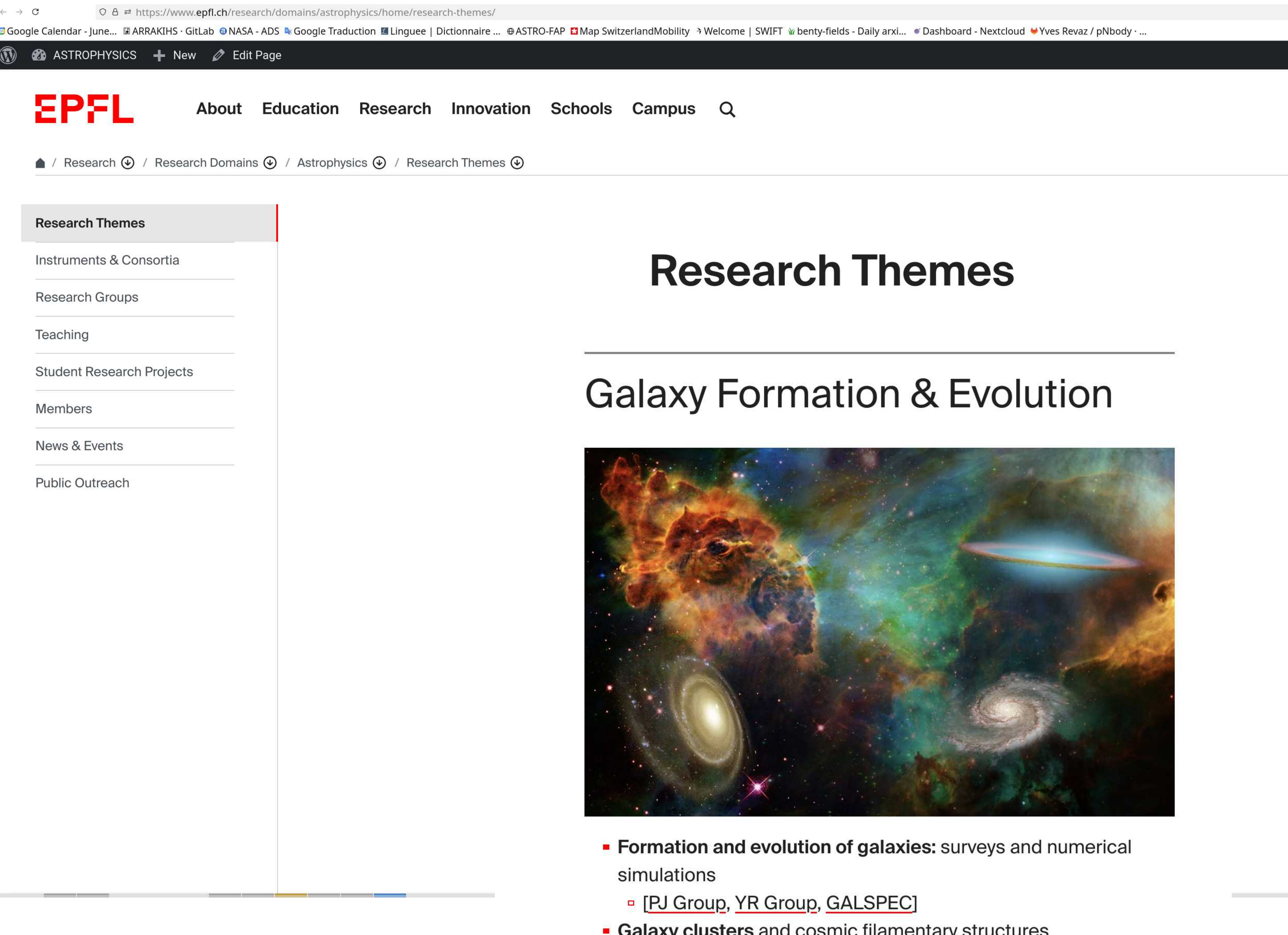
Research Methods:

Observations

Image processing

Numerical simulations

Machine learning



Research Themes

Galaxy Formation & Evolution

- **Formation and evolution of galaxies:** surveys and numerical simulations
    - [PJ Group, YR Group, GALSPEC]
- **Galaxy clusters** and cosmic filamentary structures

Introduction

Outlines of the 14th lectures

Introduction

Goal of the course

Teach how a system (stellar or galactic) evolves under the forces of gravity that are generated by itself

Ξ

Evolution of a self-gravitating system

Introduction

Outlines

Week 1:

Introduction

- The standard model in cosmology
- The formation of the large scale structure
- Which physics
- Our galaxy the Milky Way
- The Local Group and beyond

Week 2:

The gravity : a long distance force

- collision-less systems : the relaxation time

Newton Mechanics (quick reminder)

The Potential Theory I

- General results
  - Newton law, gravitational field force and potential

Week 3:

The Potential Theory I

- Spherical systems
  - Newton’s theorems
  - Circular speed, circular velocity, circular frequency, escape speed, potential energy
  - Useful relations for spherical systems

Week 4:

The Potential Theory II

- Examples of spherical models:
  - “Potential based” models
  - “Density based” models
- Axisymmetric models for disk galaxies
  - “Potential based” models
  - Potential of flattened systems
  - The potential of infinite thin (razor) disks (potential of a ring)
  - Potential of ellipsoidal systems
  - Potential of infinite thin disks and slabs

Week 5:

Stellar Orbits I

- Generalities : why studying stellar orbits ?

- Lagrangian and Hamiltonian mechanics (quick reminder)
  - Euler-Lagrange equations
  - Hamilton’s equations

- Orbits in spherical potentials
  - angular momentum conservation
  - equations of motion
  - radial orbits
  - non radial orbits

- Examples
  - Keplerian orbits
  - Orbits in an homogeneous sphere
  - Orbits in isochrone potentials

Week 6:

Stellar Orbits II

- Orbits in axisymmetric potentials
  - orbits in the equatorial plane
  - orbits outside the equatorial plane
  - equations of motion
  - orbits in the meridian plane
  - examples

Week 7:

Stellar Orbits III

- Nearly circular orbits
  - Epicycle frequences
  - The Oort constants
  - Probing the mass in the stellar disk

- Surface of section
  - Integral of motions
  - Poincaré maps

Week 8:

Stellar Orbits IV

- Orbits in planar non-axisymmetric potential
    - surface of sections

- Orbits in non-axisymmetric rotating potential
    - the Jacobi integral
    - Lagrange points
    - stability of orbits around Lagrange points
    - orbits not confined to Lagrange points

- Weak bars
    - the Lindblad resonances
    - orbit families in realistic bars

Week 9:

Equilibria of collisionless systems I

- The collisonless Boltzmann equation
  - The distribution function (DF) of stellar systems
  - The Collisionless Boltzmann equation
  - Limitations
- Relations between DFs and observables
  - Density, velocity distribution function, mean velocity, velocity dispersion
- The Jeans theorems
  - Solutions of the Collisionless Boltzmann equation
  - Symmetry and integrals of motion

Week 10:

Equilibria of collisionless systems II

- Self-consitent spherical models with Ergodic DF
  - DFs from mass distribution
    - The Eddington formula
    - Examples
  - Models defined from DFs
    - Polytropes and Plummer models
    - Parallel with hydrostatics polytropes
    - Isothermal models
    - Parallel with hydrostatics isothermal models

Week 11:

Equilibria of collisionless systems III

- Anisotropic distribution function in spherical systems
  - Motivations
  - General concepts
  - Example of an anisotropic DF
  - Application to the Hernquist model

- The Jeans Equations (moments equation)
  - Motivations
  - The Jeans Equations and conservation laws
  - The Jeans Equations in Spherical and Cylindrical coordinates

Week 12:

Equilibria of collisionless systems IV

- The Virial Equation and Virial Theorem
    - Theory
    - Applications

Stability of collisionless systems I

- Nbody- experiments
    - Are systems defined from a DF that solve the CB stable ?

Week 13:

Stability of collisionless systems II

- Linear response theory
    - in fluid systems
    - in stellar systems

- The Jeans instability
- The stability of uniformly rotating systems

Week 14:

Stability of collisionless systems III

- The stability of rotating disks : spiral stuctures
  - Spirals properties
  - The dispersion relation for a razor thin fluid disk
  - The WKB approximation

- The origin of spiral structures: another view

- Vertical instabilities
  - Nature is always more tricky...

Polycop... ? No.

- PDF manuscript notes ?
  - yes, on moodle.epfl.ch
- Recordings ?
  - No (except when I will be absent...)
- Additional material ?
  - yes, on moodle.epfl.ch

Is it a difficult course ?

- Theoretical lecture (a lots of equations)

Physics:

- Newtonwian gravity 👍
- Lagrangian/Hamiltonian formalism
- Distribution function
- A lots of paralell between different fields in physics:
  e.g. thermodynamics/statistical physics, hydrodynamics

Mathematics:

- Differencial equations, Fourrier transform, Abel integral, Elliptical coordinates

Integration over $d^3\vec{v}$ + Poisson

$$\bar{\rho}_{s1}(\bar{k},t) = -\frac{4\pi G}{k^2}\, i \int d^3\bar{v}\; \bar{k}\cdot\frac{\partial \rho_0}{\partial \bar{v}} \int_{-\infty}^{t} dt'\, e^{i\vec{k}\cdot\bar{v}(t'-t)} \left[\bar{\rho}_{s1}(k,t') + \bar{\rho}_e(k,t') \right]$$

In temporal Fourier space

$$\tilde{\bar{\rho}}_{s1}(\bar{k},\omega) = \int dt\; \bar{\rho}_{s1}(\bar{k},t)\, e^{-i\omega t}$$

this term may diverge if $\bar{k}\cdot\bar{v} = \omega$

$$\tilde{\bar{\rho}}_{s1}(\bar{k},\omega) = \left(-\frac{4\pi G}{k^2} \int \frac{d^3\bar{v}}{\bar{k}\cdot\bar{v}-\omega}\; \bar{k}\cdot\frac{\partial \rho_0}{\partial \bar{v}} \right) \left(\tilde{\bar{\rho}}_{s1}(\bar{k},\omega) + \hat{\bar{\rho}}_e(\vec{k},\omega) \right)$$

In absense of perturbation

$\rho_e = 0$

we must have :

$$\boxed{-\frac{4\pi G}{k^2} \int \frac{d^3\bar{v}}{\bar{k}\cdot\bar{v}-\omega}\; \bar{k}\cdot\frac{\partial \rho_0}{\partial \bar{v}} = 1}$$

$\mathrm{Im}(\omega) > 0$

(instead, we may have a divergence)

This is our dispertion relation

$$\omega = \omega(\vec{k}, \rho_0)$$

Exam

- **Oral Exam:**
  - Classical form : general questions on the lectures

Bibliography

- James Binney & Scott Tremaine
  - Galactic Dynamics, 2nd edition, Princeton Series in Astrophysics, Princeton University Press, 2008
- Landau & Lifshitz
  - Mechanics, 3nd edition Volume 1, Butterworth Heinemann, 1976
- Landau & Lifshitz
  - Fluid Mechanics, 2nd edition Volume 6, Butterworth Heinemann, 1987
- Landau & Lifshitz
  - Statistical Physics, 3rd edition Part 1, Volume 5, Butterworth Heinemann, 1980
- N. Deruelle & J.-P. Uzan
  - Théories de la Relativité, Belin, 2015
- S. Chandrasekhar
  - An Introduction to the Study of Stellar Structure, Dover Publications, 1939
- S. Chandrasekhar
  - Principles of Stellar Dynamics, Dover Publications, 1942
- K. F. Ogorodnikov
  - Dynamics of Stellar Systems, Pergamon Press, 1965
- D. Mihalas, B. Weibel Mihalas
  - Fundation of Radiation Hydrodynamics, Oxford University Press, 1984
- J. Binney, J. Kormendy & S.D.M. White
  - Morphology and Dynamics of Galaxies, Saas-Fee Advanced Course #3

Acknowledgements

- Daniel Pfenniger
- Pierre North
- George Meylan
- Jean-Paul Kneib

Introduction

The standard model in cosmology

a quick overview

The standard model in cosmology

- Einstein equations (1915)

$$\boxed{R_{\mu\nu} - \frac{1}{2} R g_{\mu\nu} = 8\pi G T_{\mu\nu}}$$

geometry of
spacetime

mass/energy
content

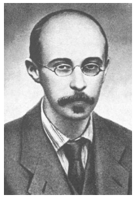
- <u>The cosmological principle:</u>

The spatial distribution of matter in the universe is homogeneous and isotropic when viewed on a large enough scale.

The standard model in cosmology

- Under the cosmological principle, Einstein equations becomes quite simple:

  The Friedman-Lemaître equations:



Alexander Friedman
(1888-1925)

Georges Lemaître
(1894-1966)

- Mass/energy content of the Universe:

$$\boxed{\Omega_{\mathrm{M}} + \Omega_{\mathrm{K}} + \Omega_{\Lambda} = 1}$$

Fraction of matter
Sensitive to the gravity

curvature

cosmological constant
or
dark energy

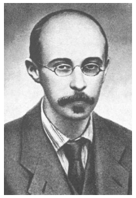
- Time evolution of its size:

$$\boxed{a(t) = a(\Omega_{\mathrm{M}}, \Omega_{\mathrm{K}}, \Omega_{\Lambda})}$$

The standard model in cosmology

- Under the cosmological principle, Einstein equations become quite simple:

  The Friedman-Lemaître equations:



Alexander Friedman
(1888-1925)

Georges Lemaître
(1894-1966)

- Mass/energy content of the Universe:

$$\boxed{\Omega_{\mathrm{M}} + \Omega_{\mathrm{K}} + \Omega_{\Lambda} = 1}$$

Fraction of matter
Sensitive to the gravity

curvature

cosmological constant
or
dark energy

- Time evolution of its size:

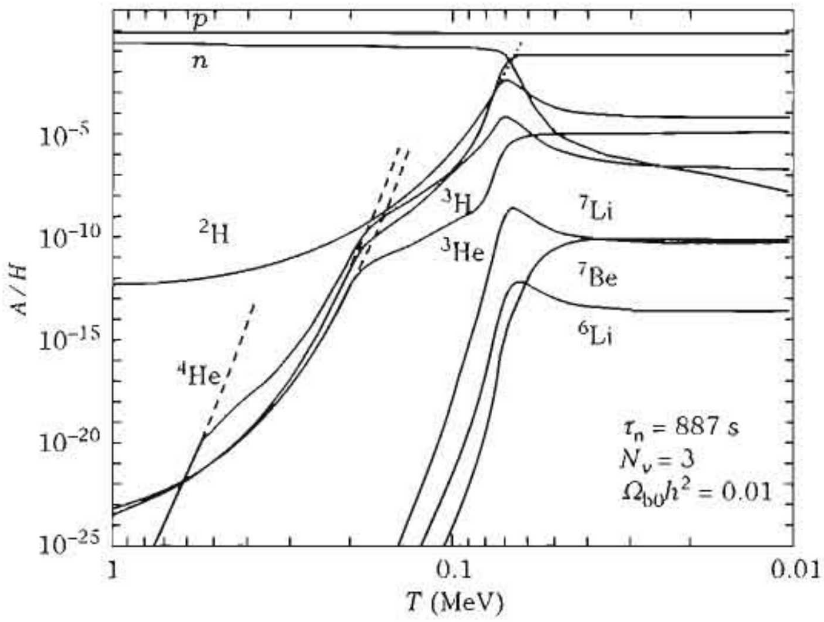
$$\boxed{a(t) = a(\Omega_{\mathrm{M}}, \Omega_{\mathrm{K}}, \Omega_{\Lambda})}$$

The standard model in cosmology

Observational constraints:

Big Bang nucleosynthesis (BBN)

Formation of lights atomic elements,
(H, He, Li, Be) during 0.1 and 200s after the BB,
when the Universe was hot and dense.



→ constraints on the fraction of atoms (baryons) in the Universe

$$\Omega_b = 0.042$$

Ralph Alpher
1948

Hans Bethe

George Gamow

Nobel Prize in Physics 1967

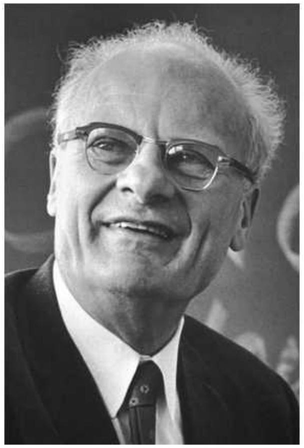
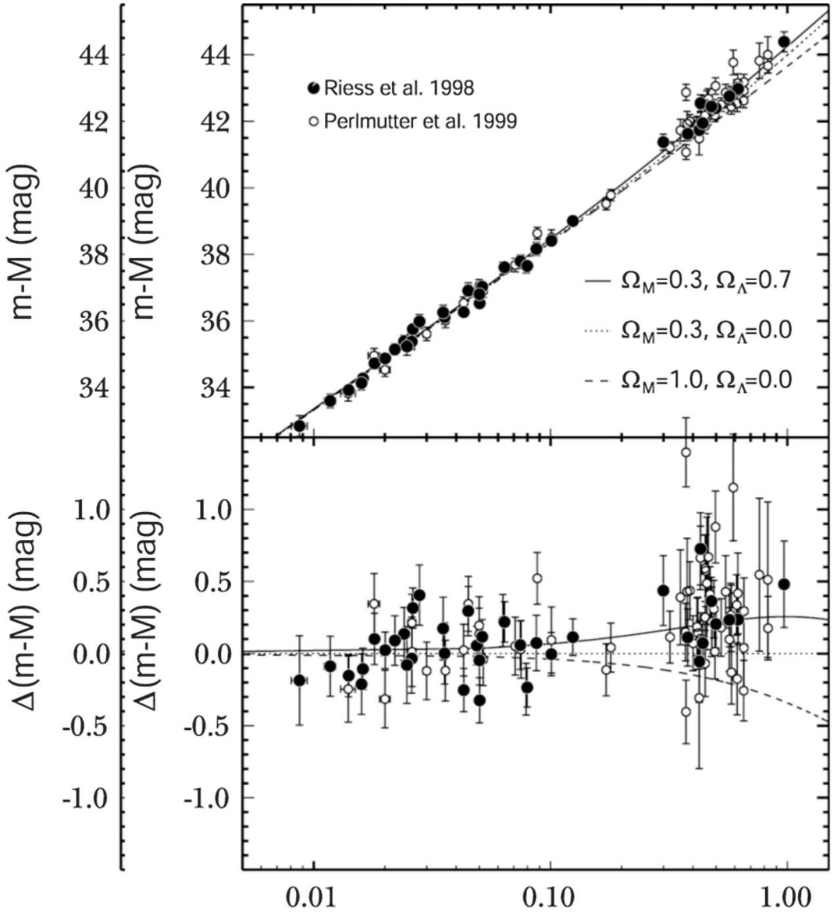


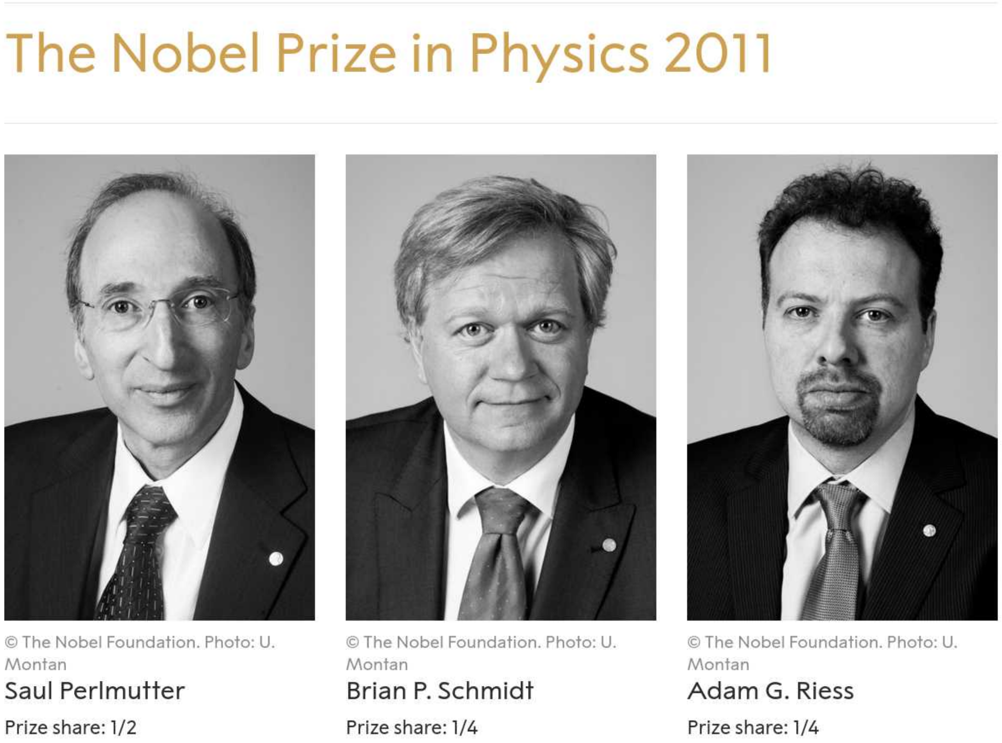
Photo from the Nobel Foundation archive.
Hans Albrecht Bethe

The standard model in cosmology

Observational constraints:

The standard candels (Type Ia supernovae)





The Nobel Prize in Physics 2011

© The Nobel Foundation. Photo: U. Montan

Saul Perlmutter

Prize share: 1/2

© The Nobel Foundation. Photo: U. Montan

Brian P. Schmidt

Prize share: 1/4

© The Nobel Foundation. Photo: U. Montan

Adam G. Riess

Prize share: 1/4

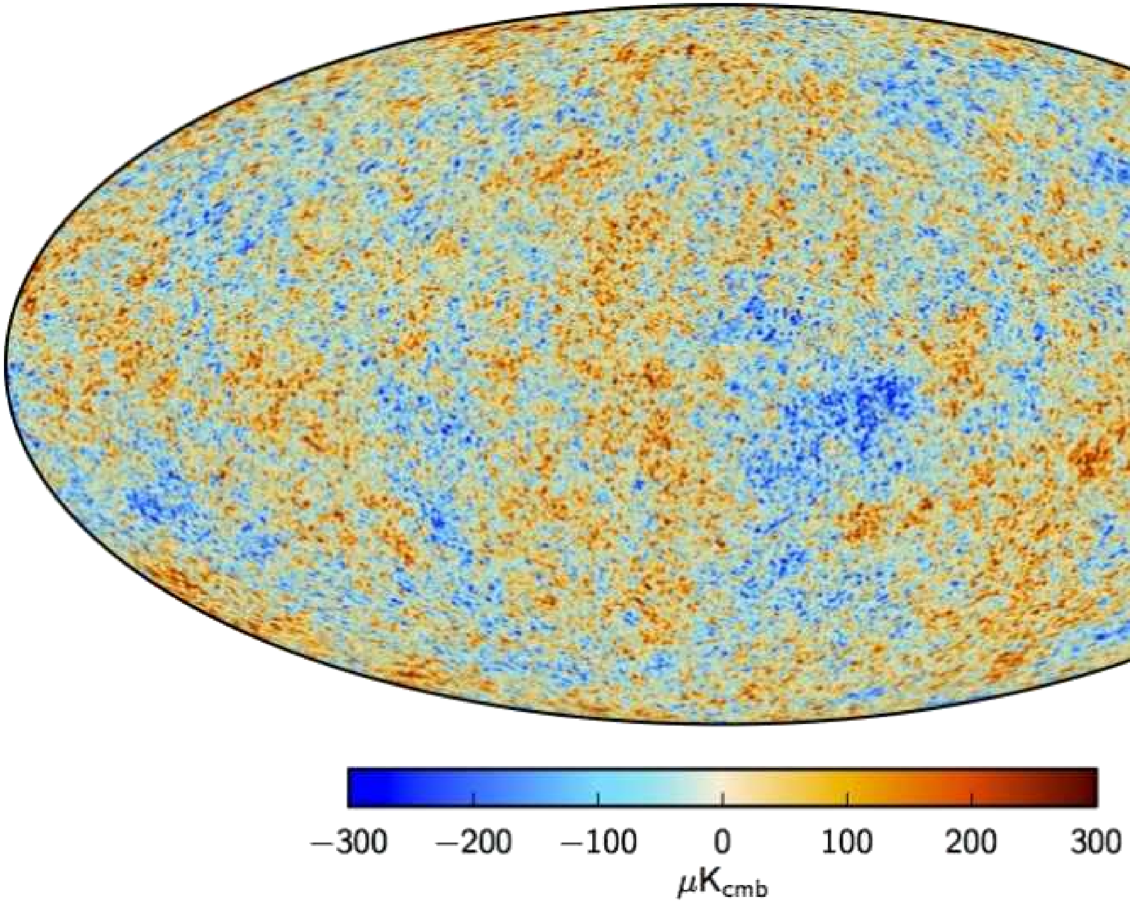
$$\Omega_\Lambda \neq 0$$

Perlmutter 1998

The standard model in cosmology

Observational constraints:

The cosmological microwave background



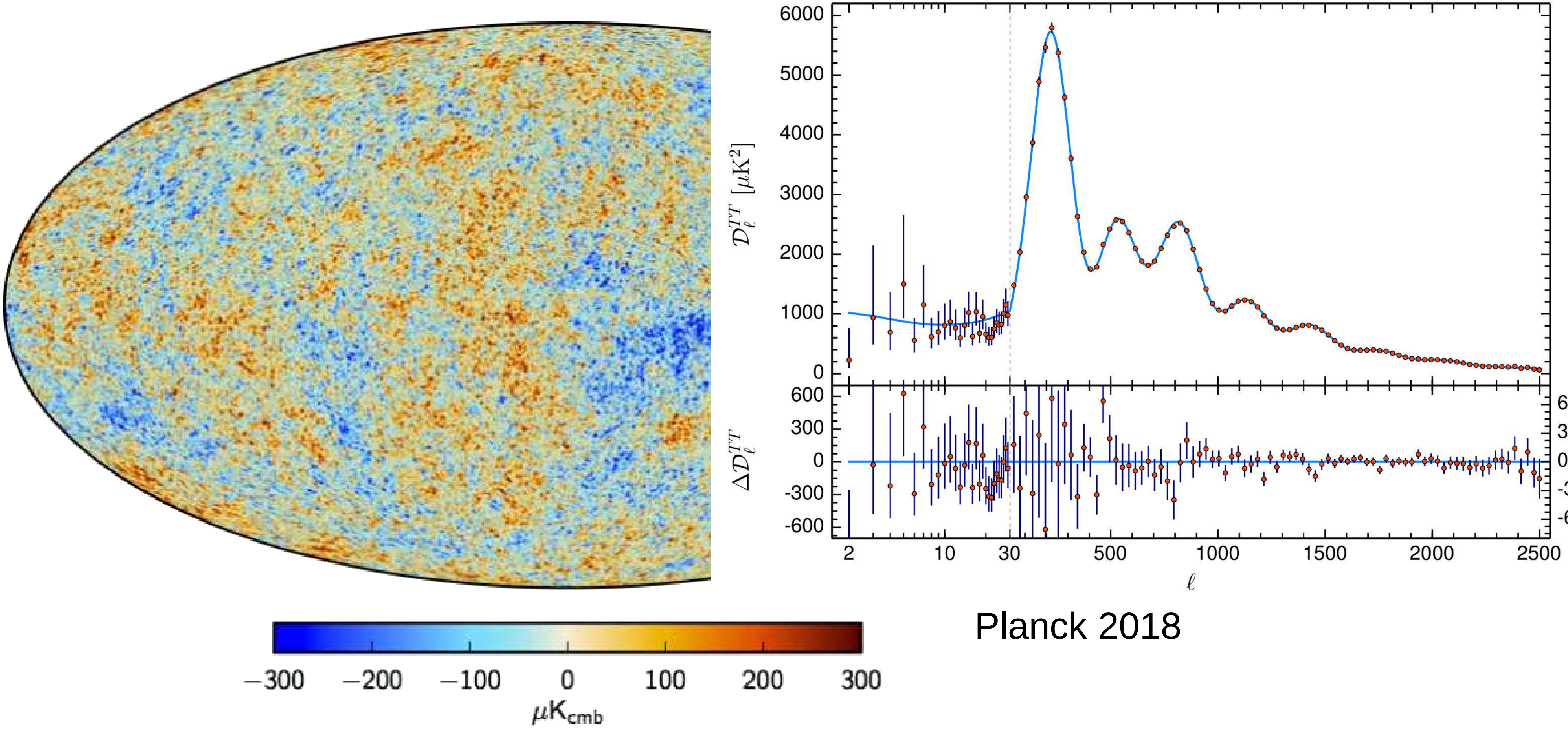
Planck 2018

$$\Omega_{\mathrm{K}} = 0 \quad \Omega_{\mathrm{M}} + \Omega_{\Lambda} = 1$$

The standard model in cosmology

Observational constraints:

The cosmological microwave background



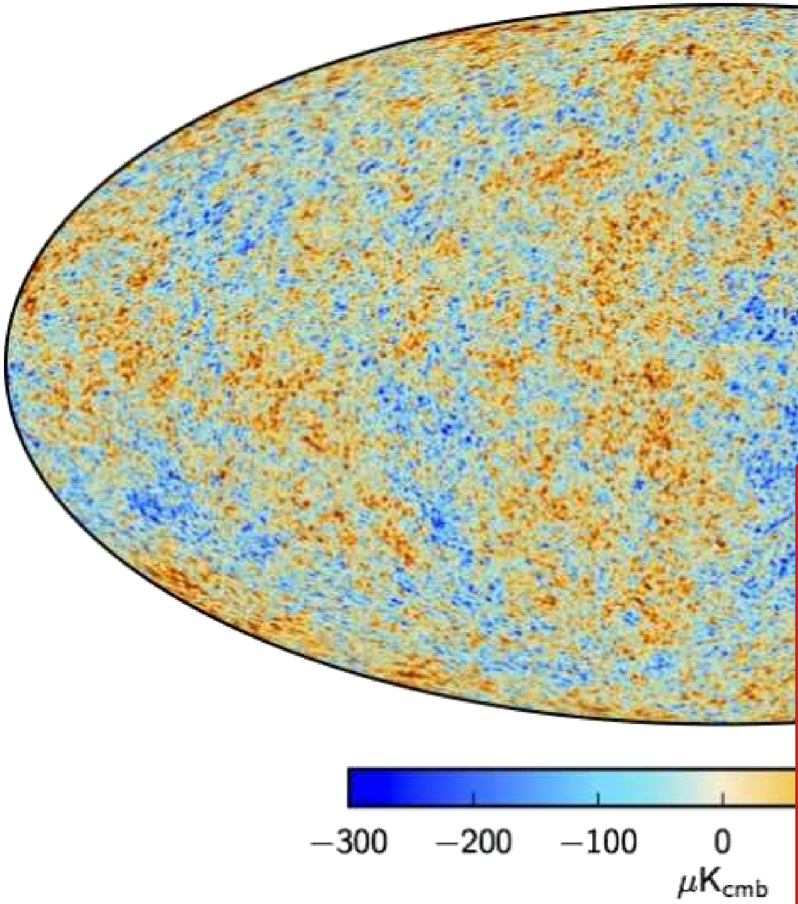
Planck 2018

$$\Omega_{\mathrm{K}} = 0 \quad \Omega_{\mathrm{M}} + \Omega_{\Lambda} = 1$$

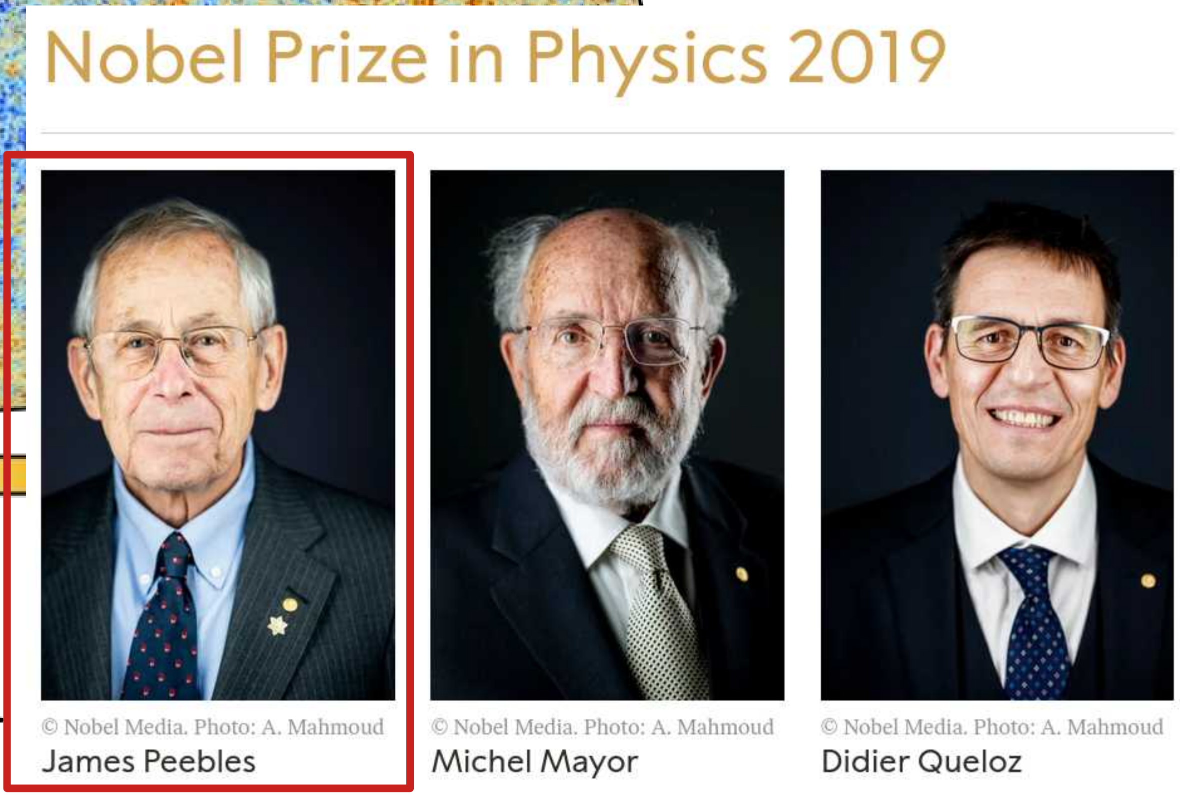
The standard model in cosmology

Observational constraints:

The cosmological microwave background



$$\Omega_K = 0 \quad \Omega_M +$$



Nobel Prize in Physics 2019

© Nobel Media. Photo: A. Mahmoud
James Peebles

© Nobel Media. Photo: A. Mahmoud
Michel Mayor

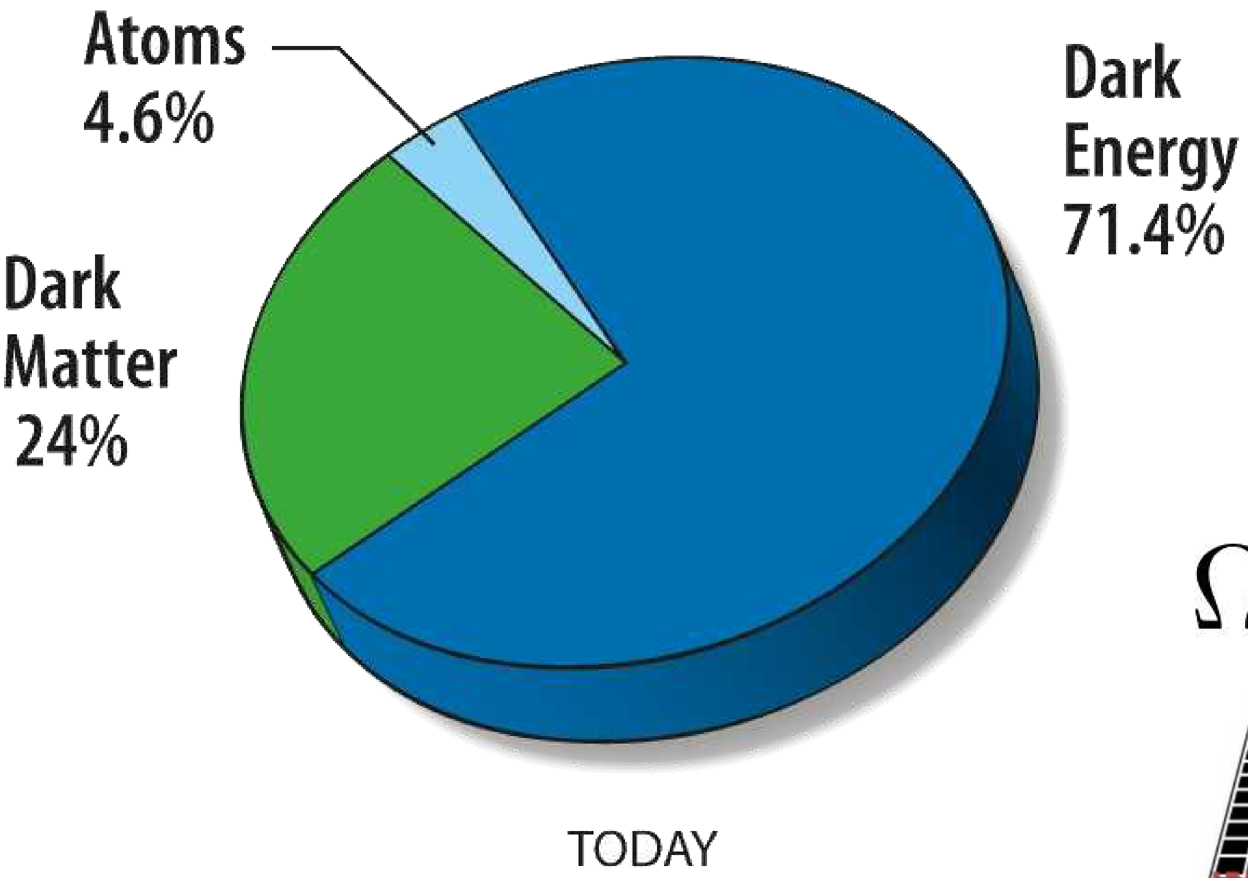
© Nobel Media. Photo: A. Mahmoud
Didier Queloz

The standard model in cosmology

ΛCDM

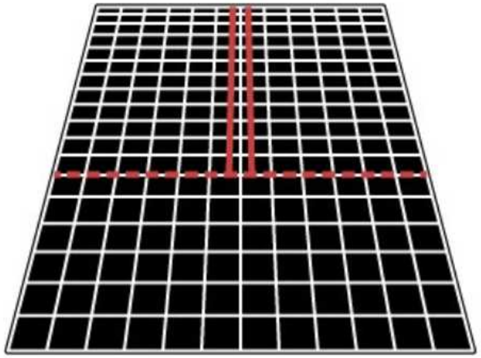
Ω_b

Ω_DM



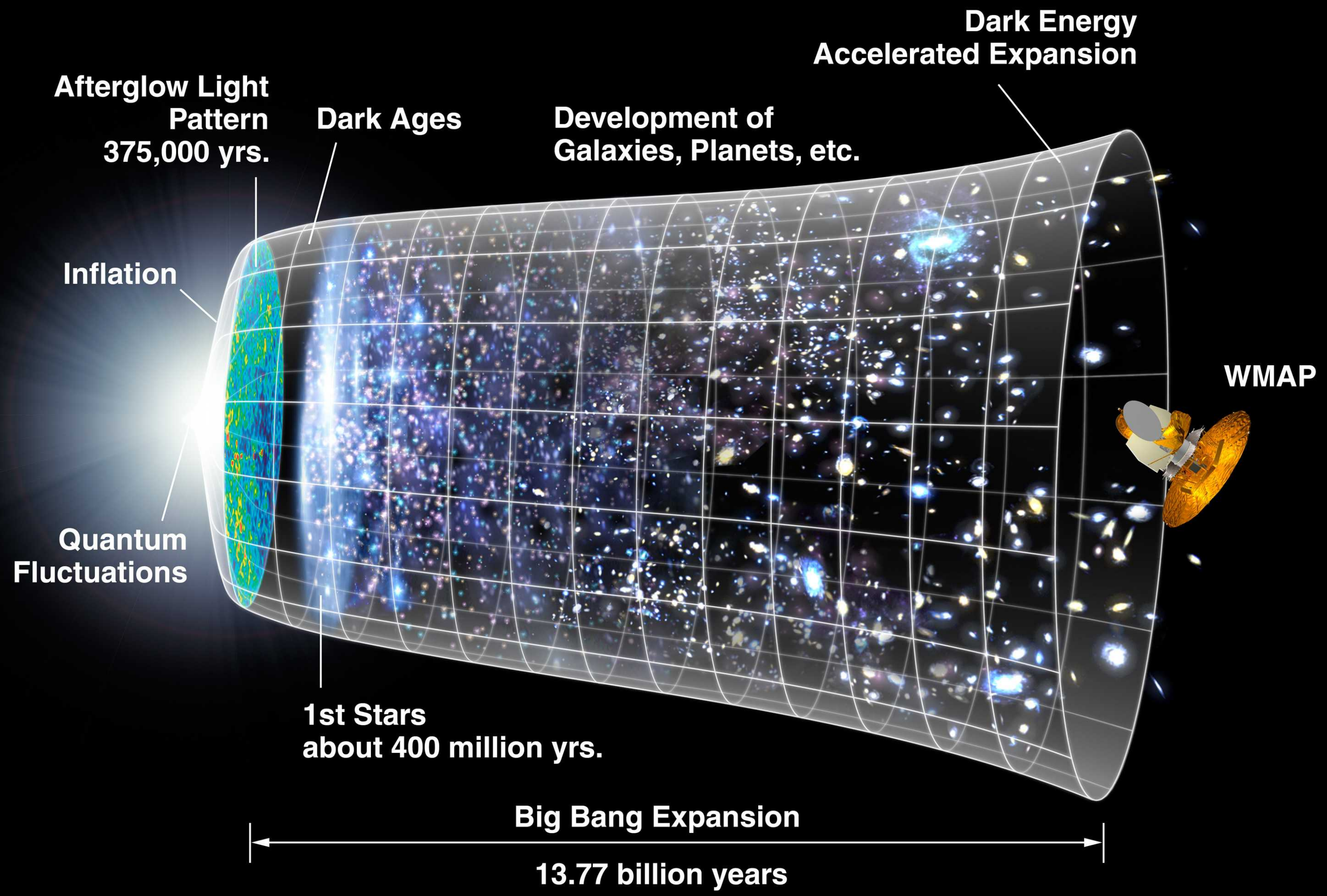
Ω_Λ

$$\Omega_\mathrm{K} = 0$$



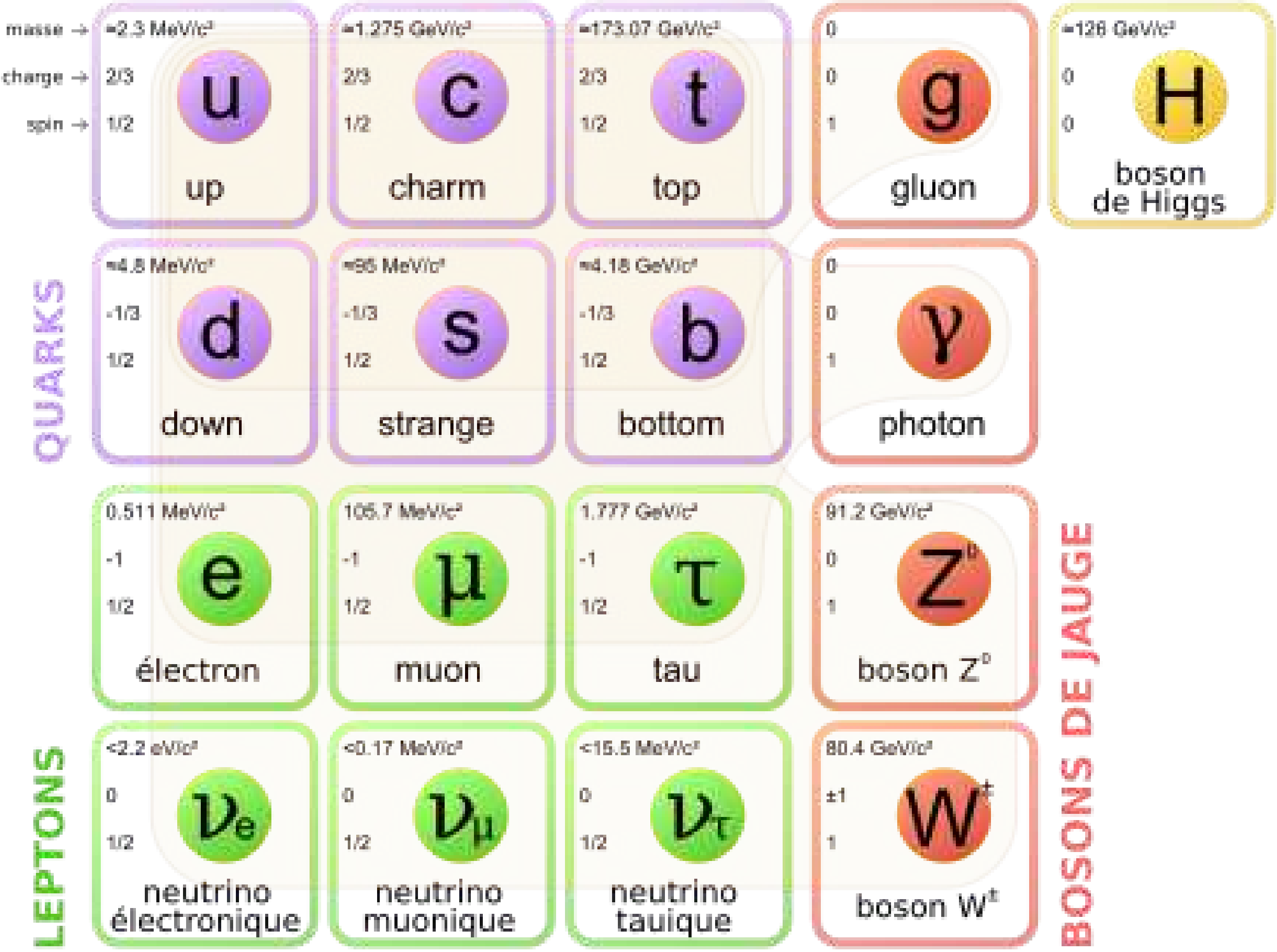
$$\Omega_\mathrm{M} = \Omega_\mathrm{DM} + \Omega_\mathrm{b}$$

$$a(t)=a(t,\Omega_M,\Omega_K,\Omega_\Lambda)$$

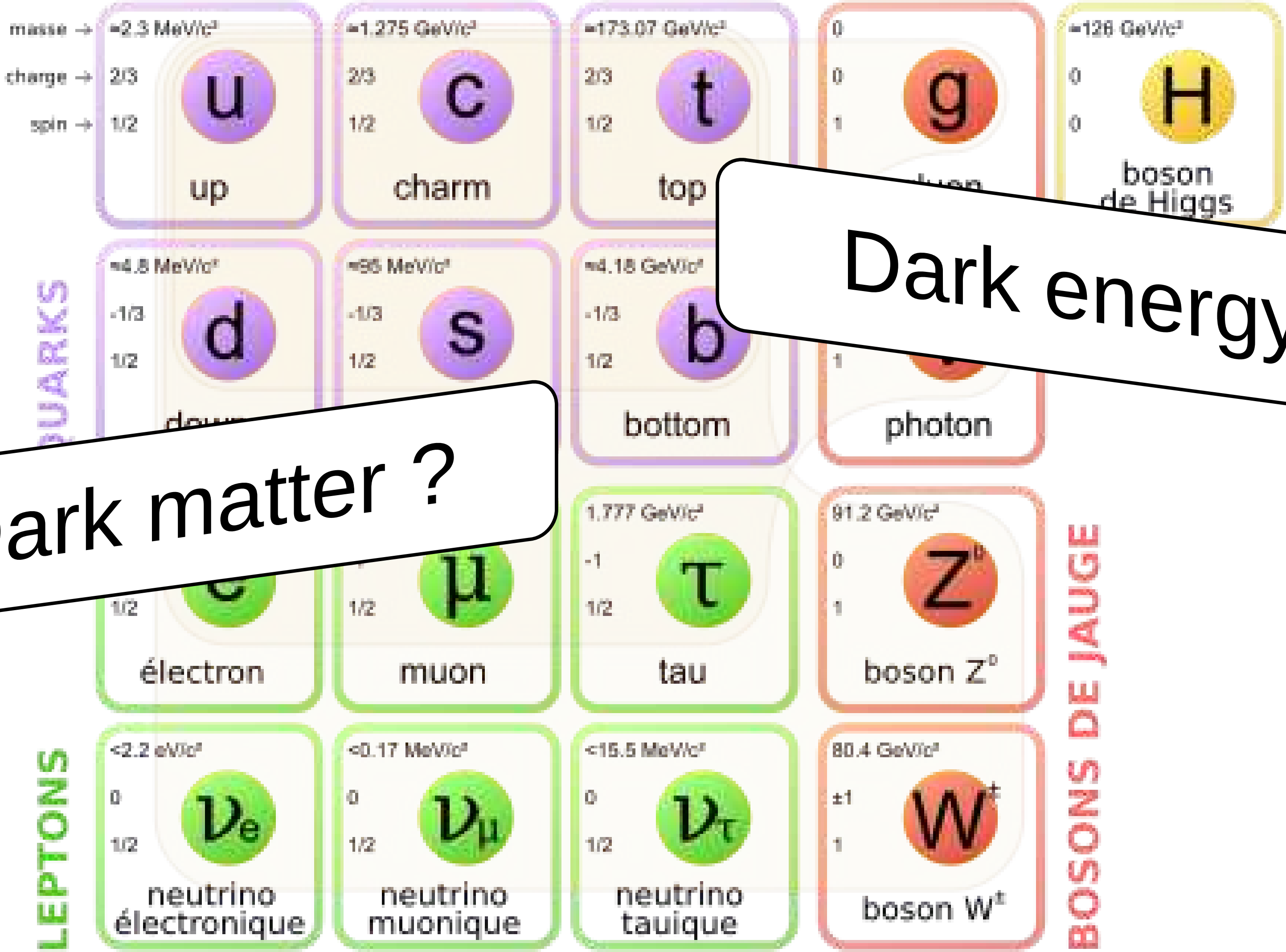


NASA/WMAP Science Team

The standard model particle physics



The standard model particle physics



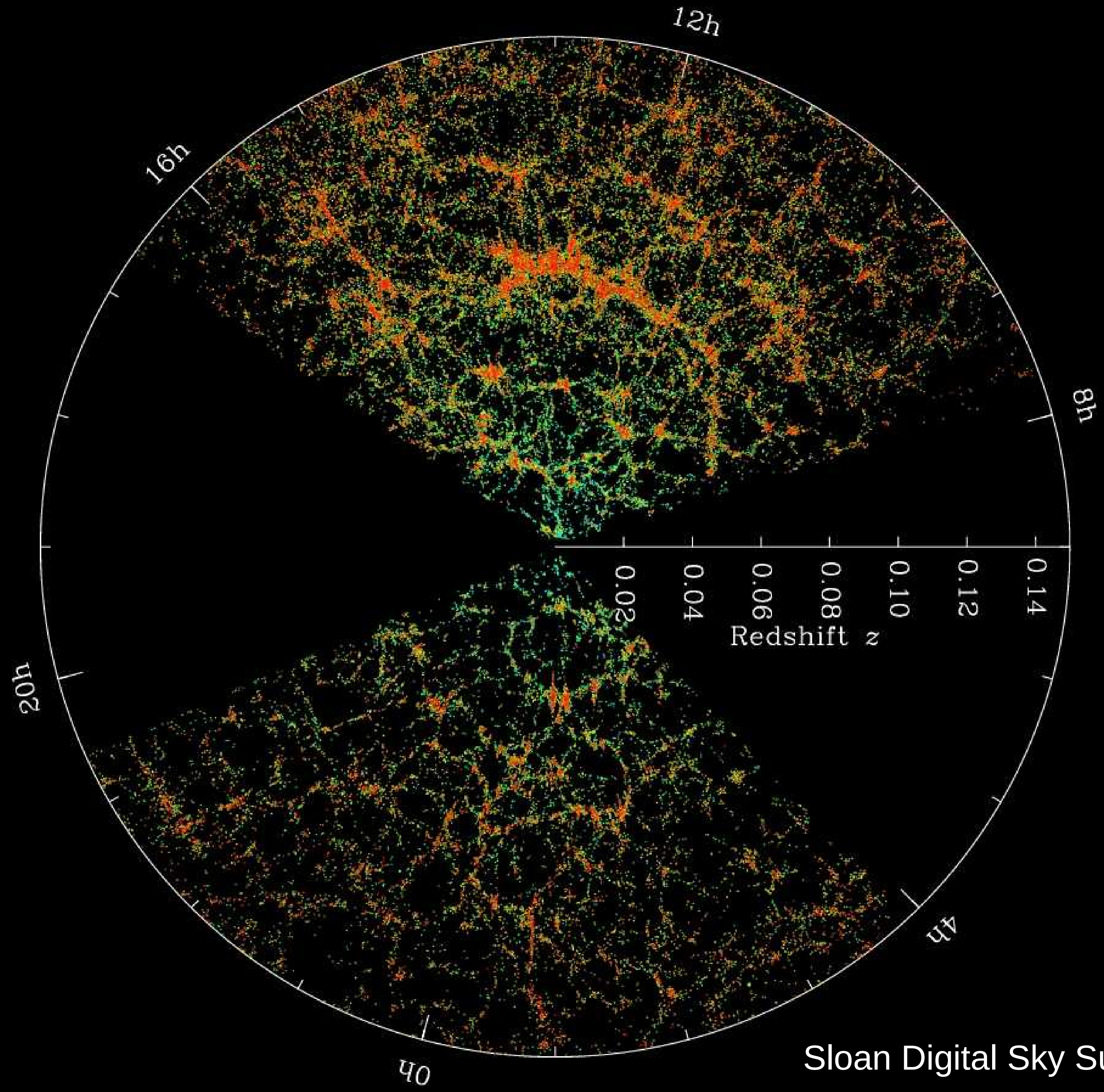
Dark energy ?

Dark matter ?

Introduction

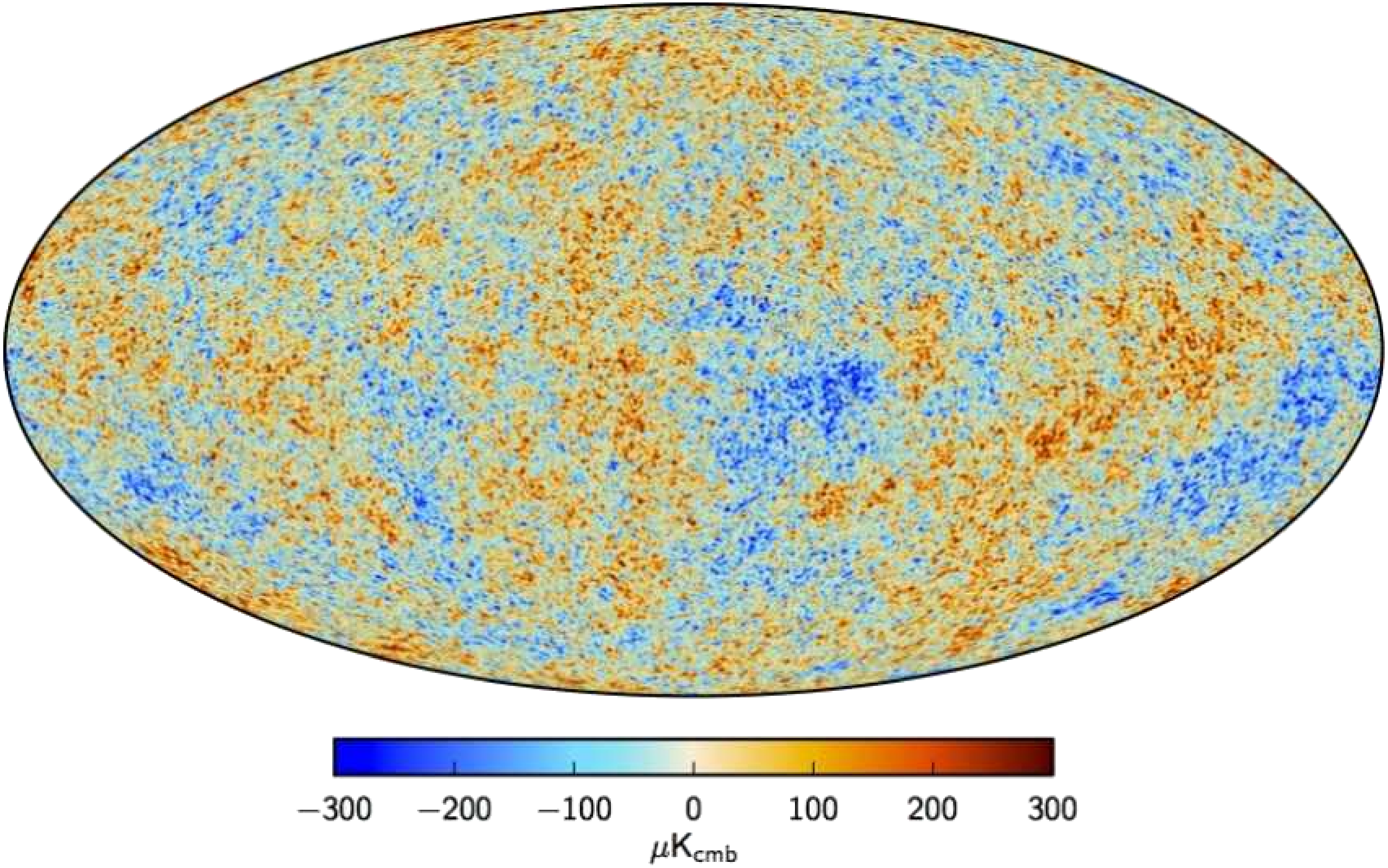
The formation of the large scale structures

a quick overview



Sloan Digital Sky Survey (SDSS)

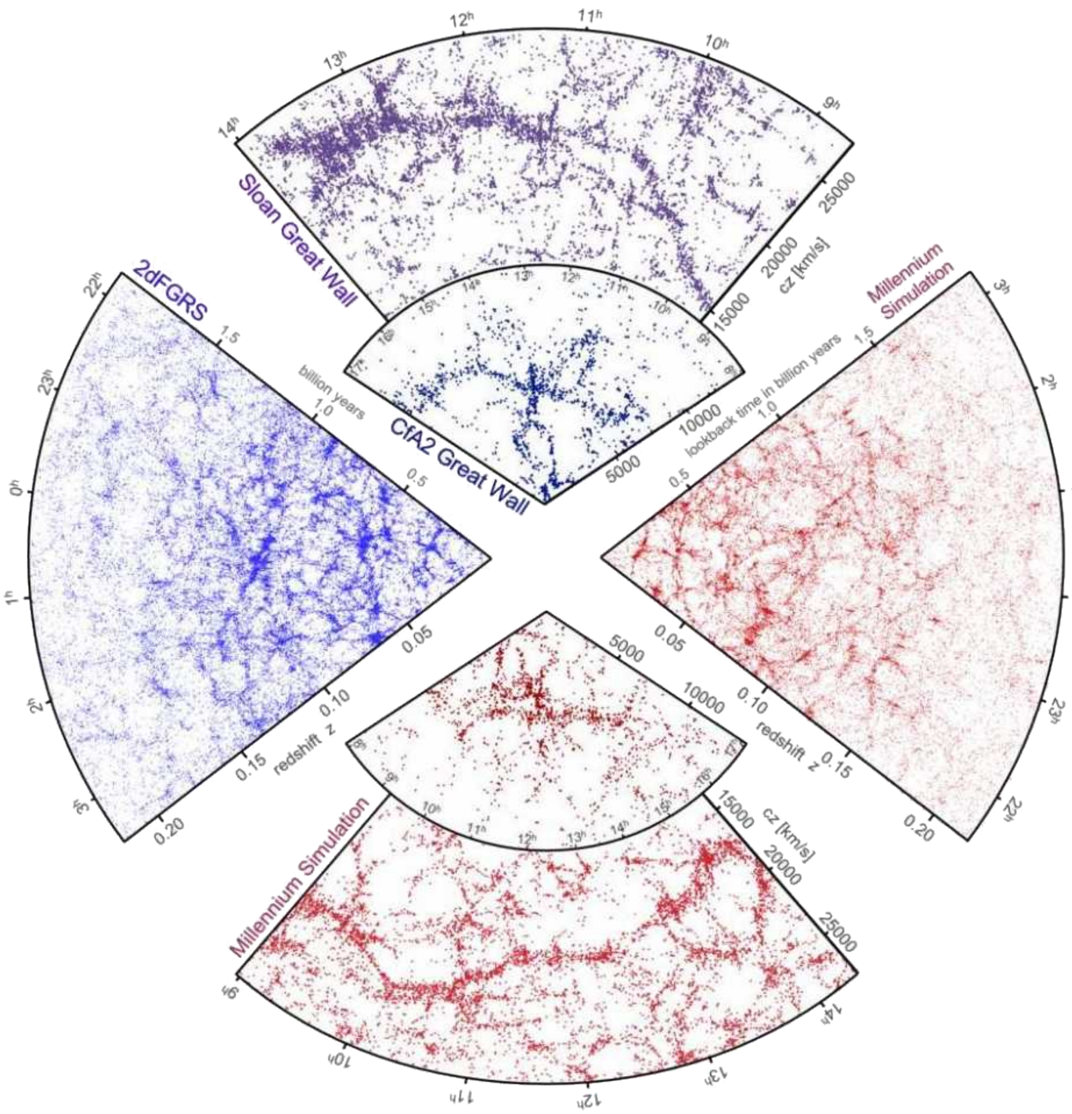
Temperature/Density fluctuations of the universe (CMB) at the recombination epoch, when it was only 380'000 years old



Credit : the Planck collaboration

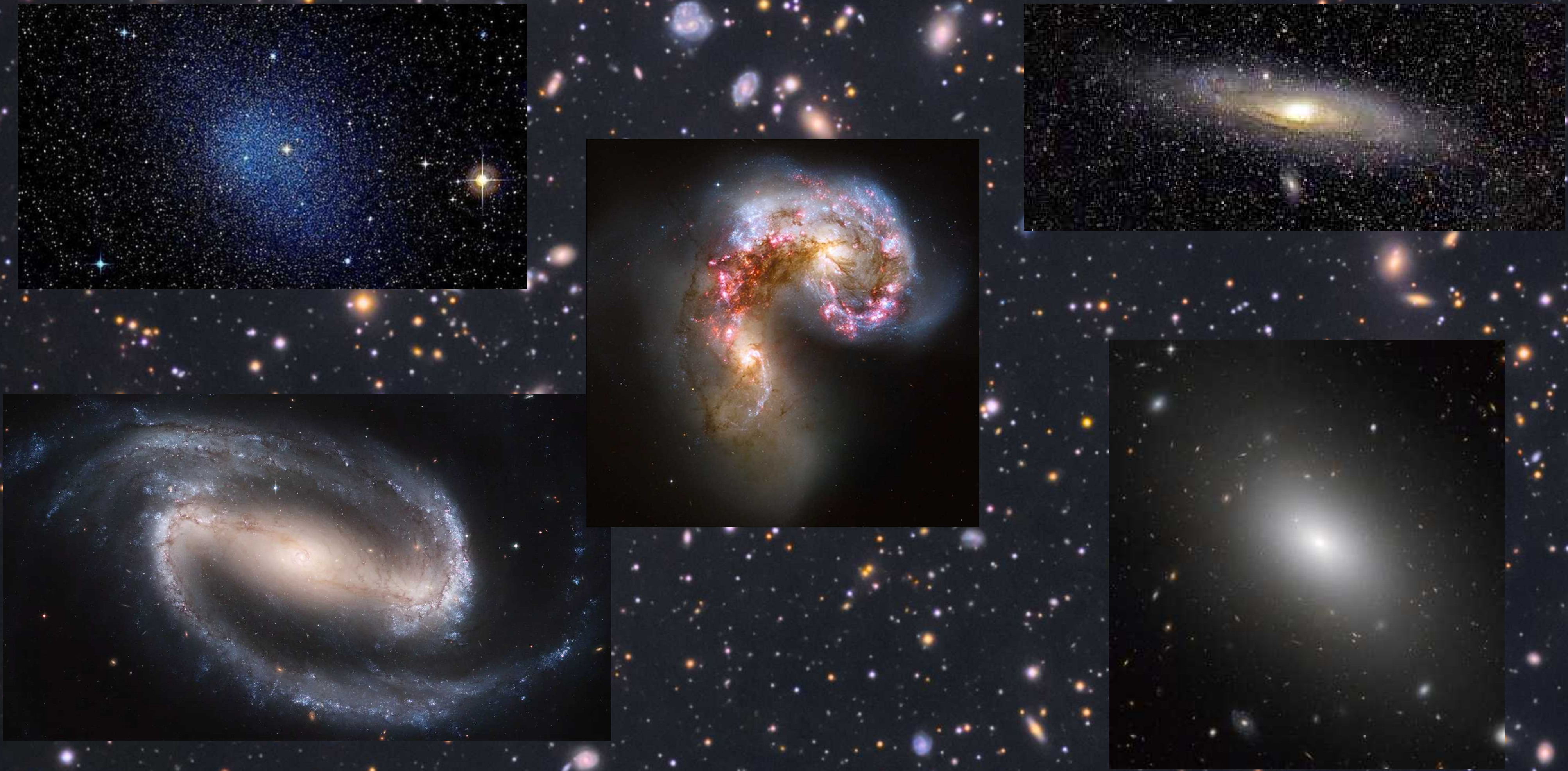
Large scale (~200 Mpc) Dark Matter Only (DMO) simulation

ΛCDM is successful at reproducing the large scale structure of our Universe



Springel et al. 2006

Reproducing and understanding the Universe at small scale, at the scale of galaxies, is much more challenging...



Introduction

Galaxy formation: Which physics ?

Galaxy formation
Which physics ?

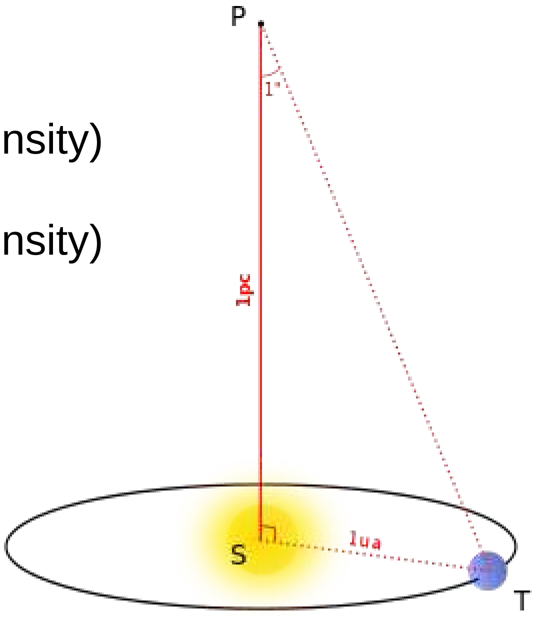
- Gravity
- Gas hydrodynamics
- Gas radiative cooling, gas heating
- Star formation
- Stellar feedback (Supernovae Ia/II, AGB, etc.)
- Chemical evolution, gas mixing, diffusion
- Active Galactic Nuclei (AGN) feedback
- Cosmic rays
- Magnetic fields
- Thermal conductivity
- Dust
- ...

Galaxy formation
Which physics ?

- Gravity
- Gas hydrodynamics
- Gas radiative cooling, gas heating
- Star formation
- Stellar feedback (Supernovae Ia/II, AGB, etc.)
- Chemical evolution, gas mixing, diffusion
- Active Galactic Nuclei (AGN) feedback
- Cosmic rays
- Magnetic fields
- Thermal conductivity
- Dust
- ...

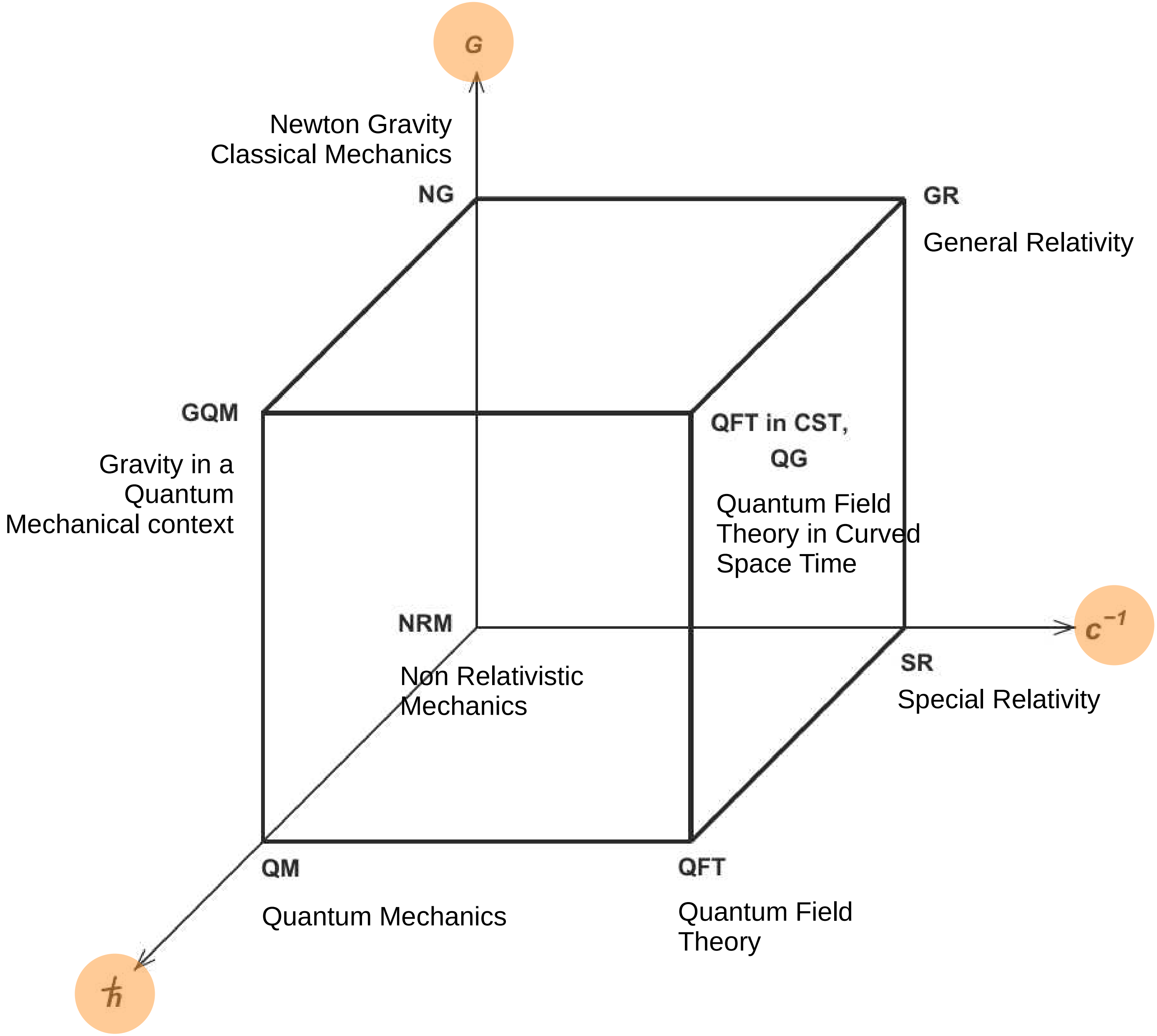
Units

Distances:	Parsec (pc)	= 3.2616 light year = 3.085x10^{16} meter
Masses:	Solar Mass ($M_\odot$)	= 2x10^{30} kg
Luminosities:	Solar Luminosity ($L_\odot$)	= 3x10^{26} Watt
Time:	Giga Year (Gyr)	= 10^9 yr
	Mega Year (Myr)	= 10^6 yr
Speed:	km/s	= km/s
Densities	atom/cm^3	= 1.7x10^{-21} kg/m^3 (air density)
	$M_\odot$ / pc^3	= 6.7x10^{-20} kg/m^3 (air density)



Credit : wikipedia

The cube of theoretical physics



Newton Gravity
Classical Mechanics

NG

GR

General Relativity

GQM

Gravity in a
Quantum
Mechanical context

QFT in CST,
QG

Quantum Field
Theory in Curved
Space Time

NRM

Non Relativistic
Mechanics

SR

Special Relativity

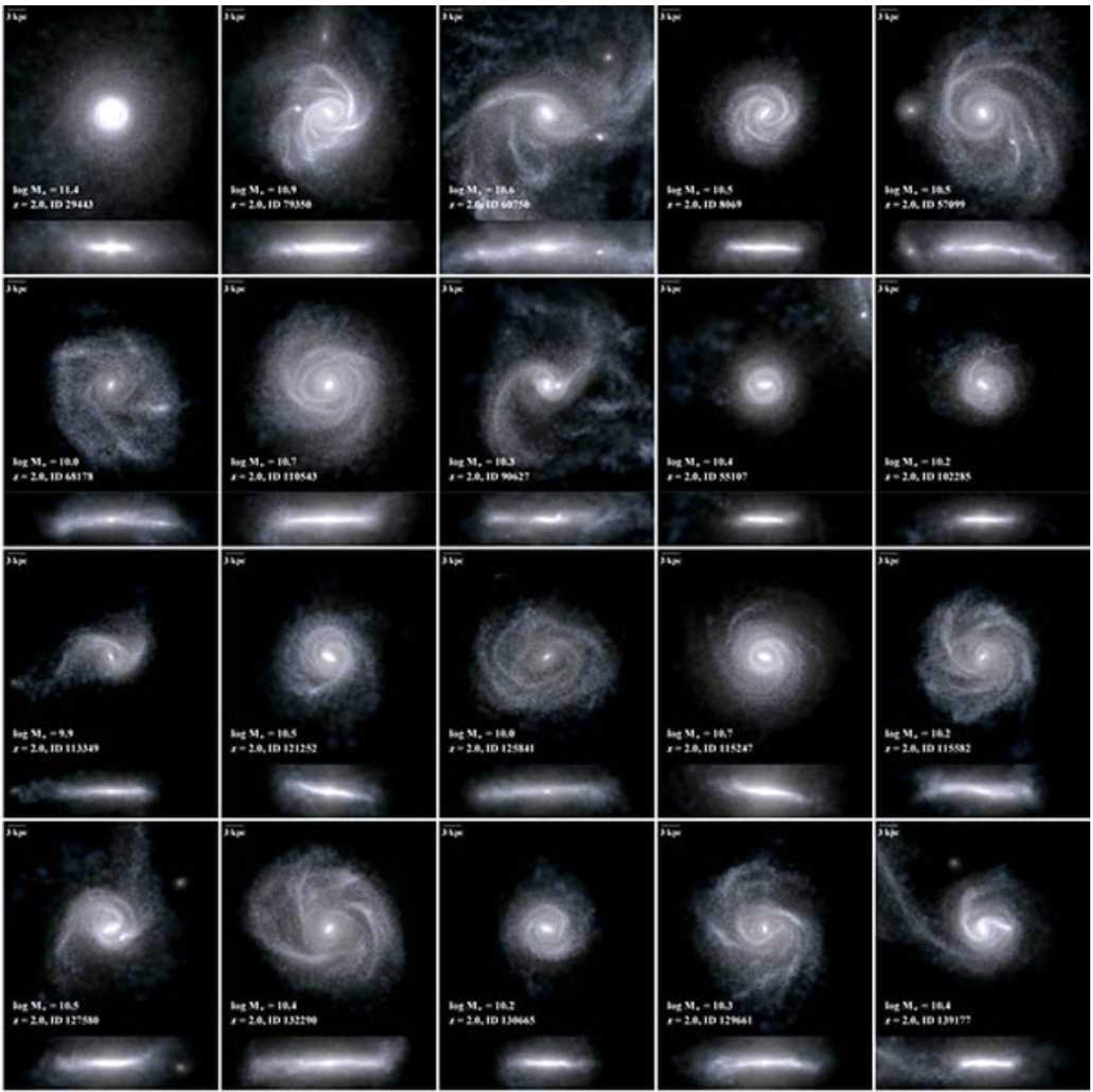
QM

Quantum Mechanics

QFT

Quantum Field
Theory

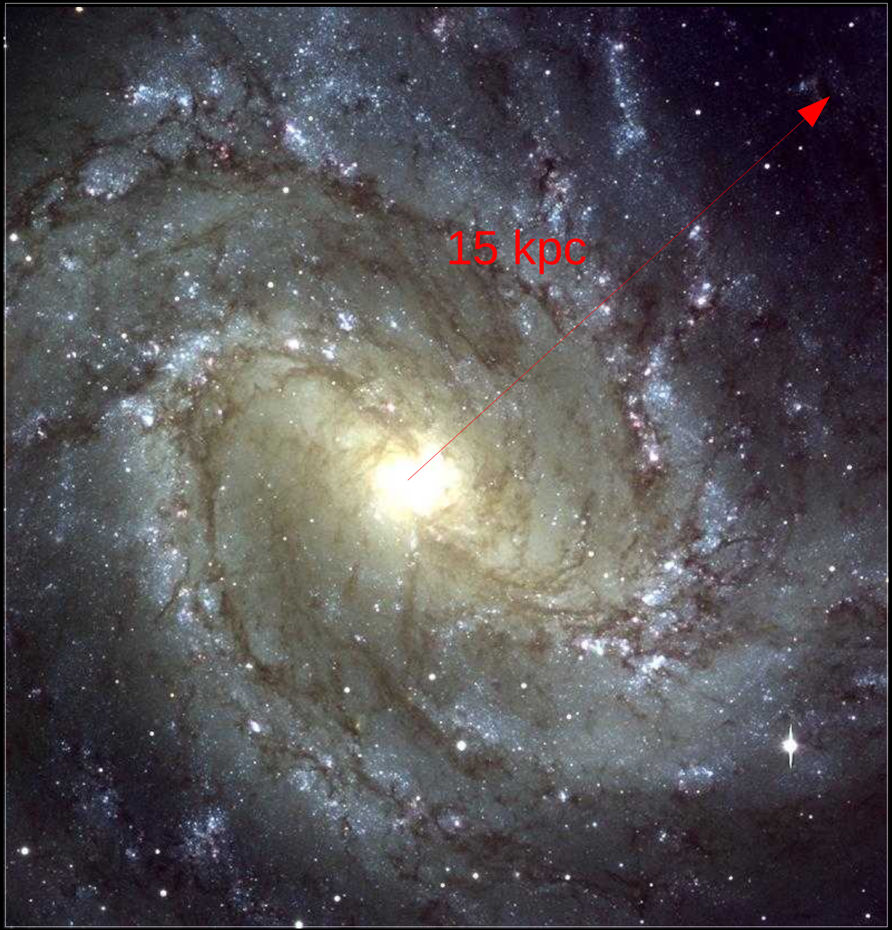
Fig. 1.1: The landscape of theoretical physics can be concisely described by a cube — The Cube of Theoretical Physics — whose axes represents the three fundamental constants $G, \hbar$ and c^{-1}. The vertices and linkages describe different structural properties of the physical theories. See text for detailed description.


3 kpc
log M_* = 11.4
z = 2.0, ID 29443
3 kpc
log M_* = 10.9
z = 2.0, ID 79350
3 kpc
log M_* = 10.6
z = 2.0, ID 60750
3 kpc
log M_* = 10.5
z = 2.0, ID 8069
3 kpc
log M_* = 10.5
z = 2.0, ID 57099
3 kpc
log M_* = 10.0
z = 2.0, ID 68178
3 kpc
log M_* = 10.7
z = 2.0, ID 110543
3 kpc
log M_* = 10.3
z = 2.0, ID 90627
3 kpc
log M_* = 10.4
z = 2.0, ID 55107
3 kpc
log M_* = 10.2
z = 2.0, ID 102285
3 kpc
log M_* = 9.9
z = 2.0, ID 113349
3 kpc
log M_* = 10.5
z = 2.0, ID 121252
3 kpc
log M_* = 10.0
z = 2.0, ID 125841
3 kpc
log M_* = 10.7
z = 2.0, ID 115247
3 kpc
log M_* = 10.2
z = 2.0, ID 115582
3 kpc
log M_* = 10.5
z = 2.0, ID 127580
3 kpc
log M_* = 10.4
z = 2.0, ID 132290
3 kpc
log M_* = 10.2
z = 2.0, ID 130665
3 kpc
log M_* = 10.3
z = 2.0, ID 129661
3 kpc
log M_* = 10.4
z = 2.0, ID 139177

Introduction

Our galaxy
The Milky Way

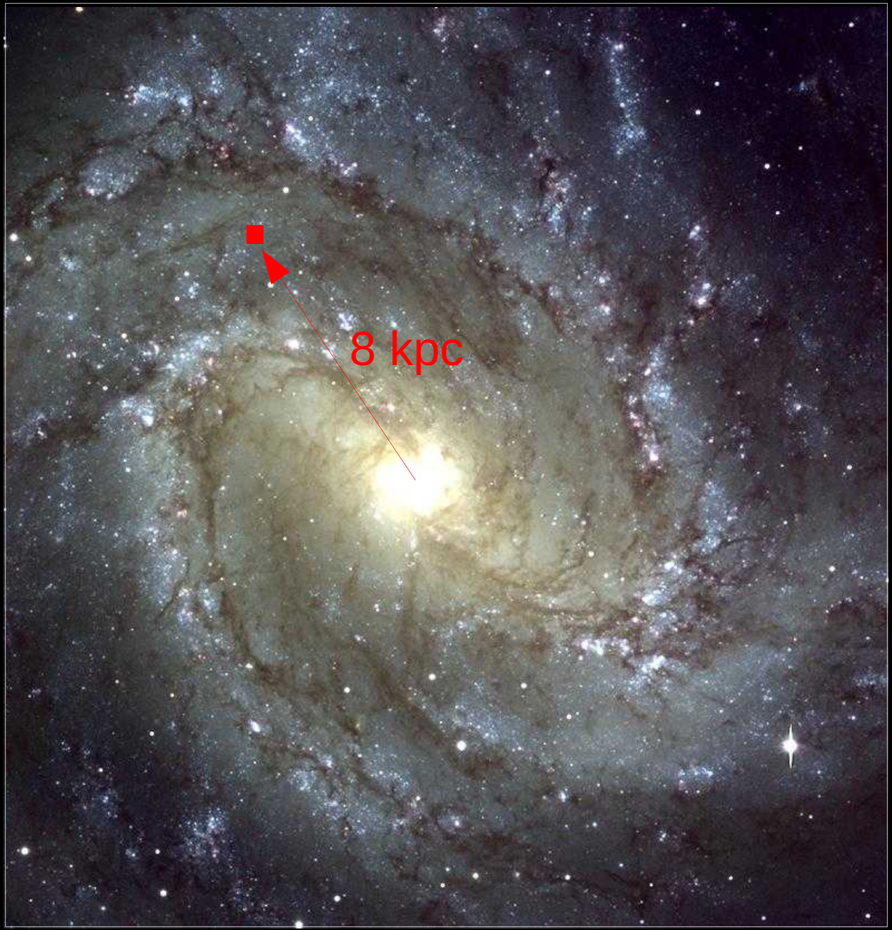
The Milky Way : a disky galaxy



M83

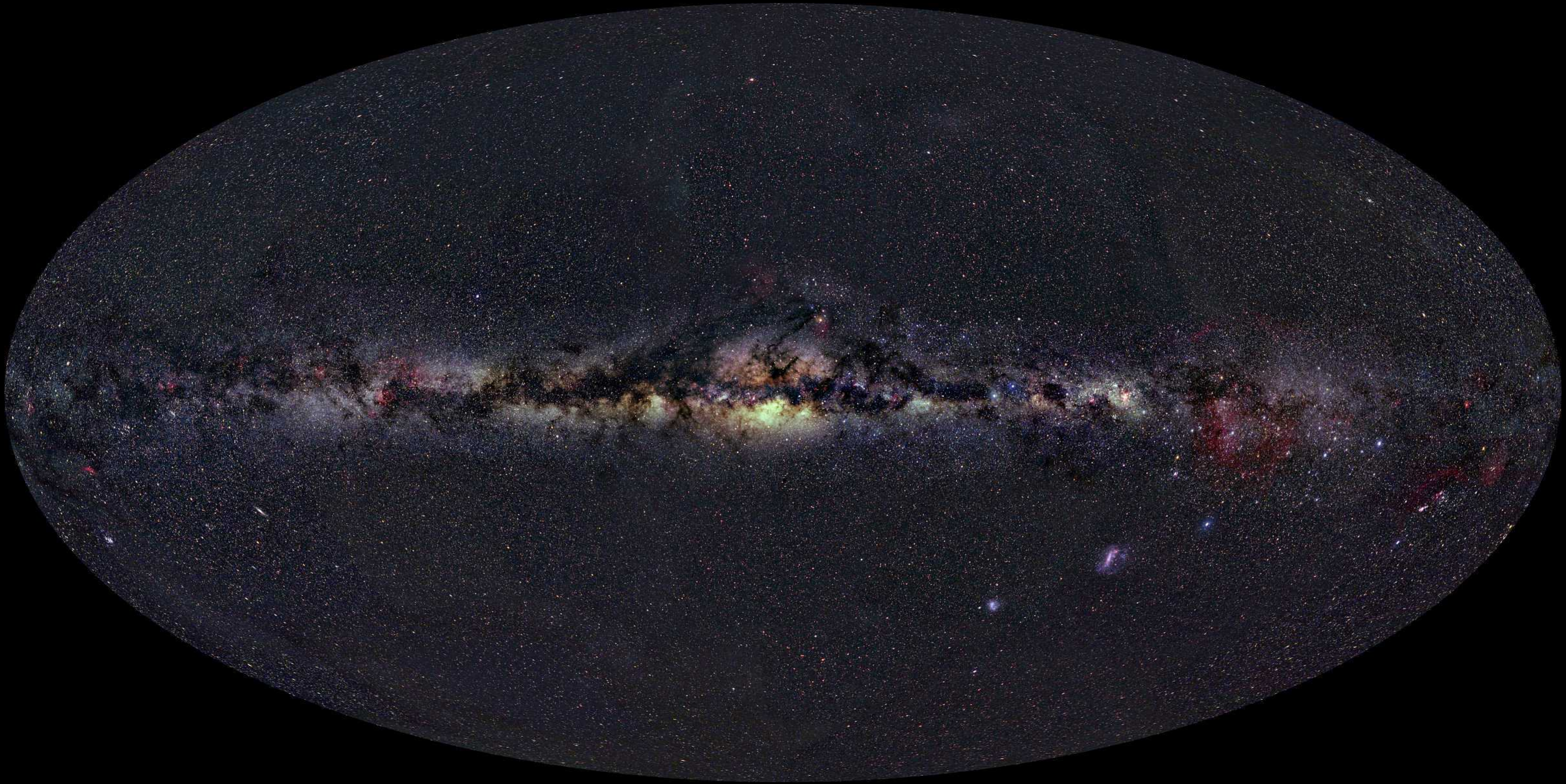
NGC4945

Position of the Sun

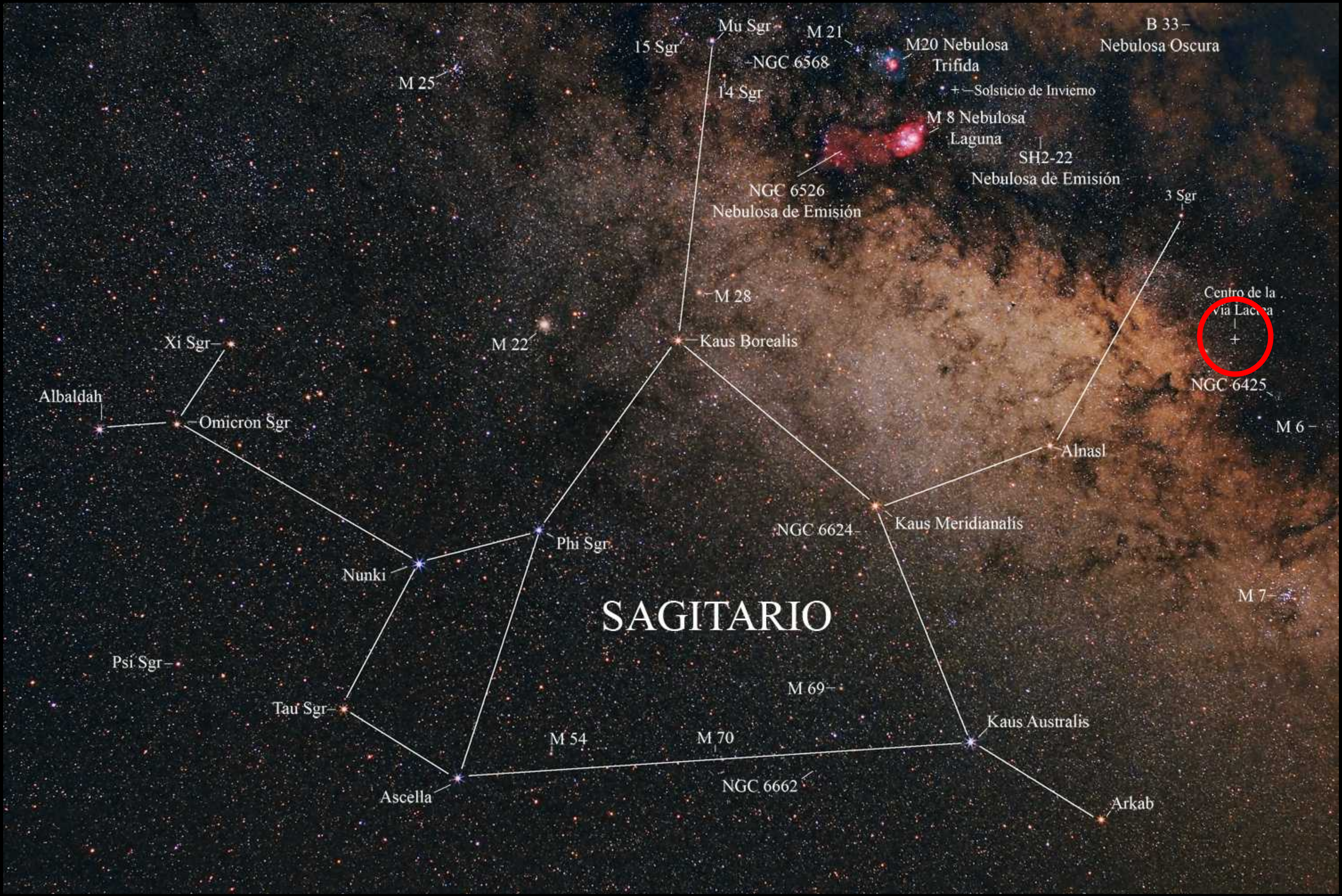


M83

NGC4945



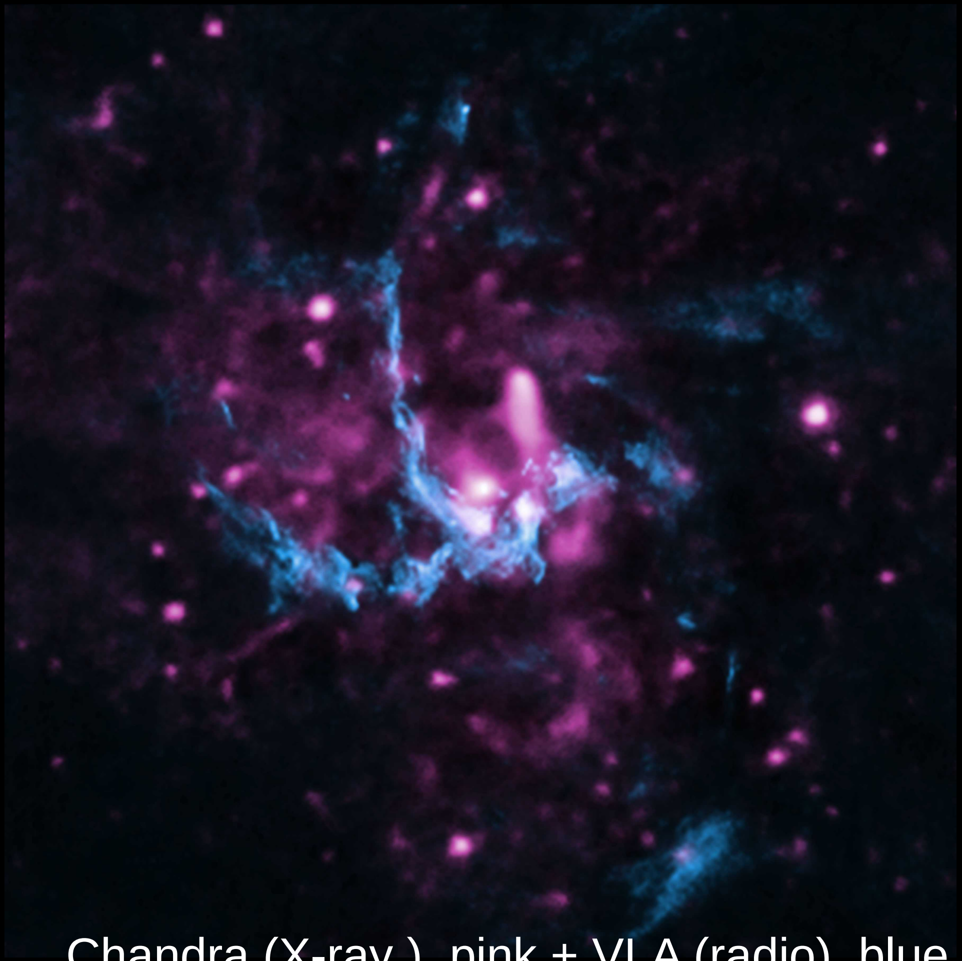
The Galactic Centre



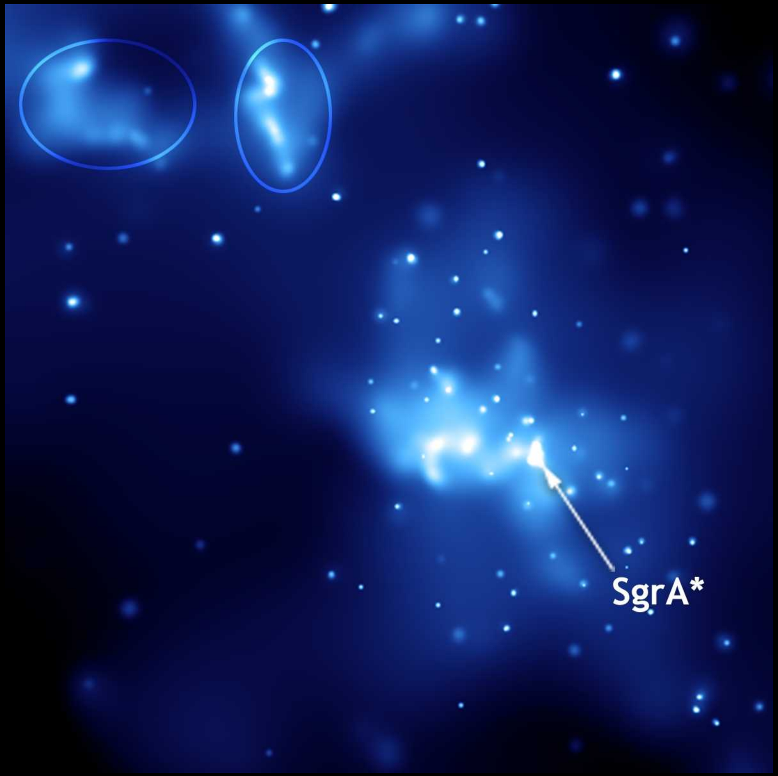
The Galactic Centre

Very well determined via radio observations of the radio-source Sagittarius A* (Galactic Black Hole)

Location : 17h45m 40.0409s (RA), -29º0’28.118’’ (DEC)
Distance: 25.900±1.400 light years (7.940±420 pc)
Mass: 4.31±0.38 10^6 $M_\odot$

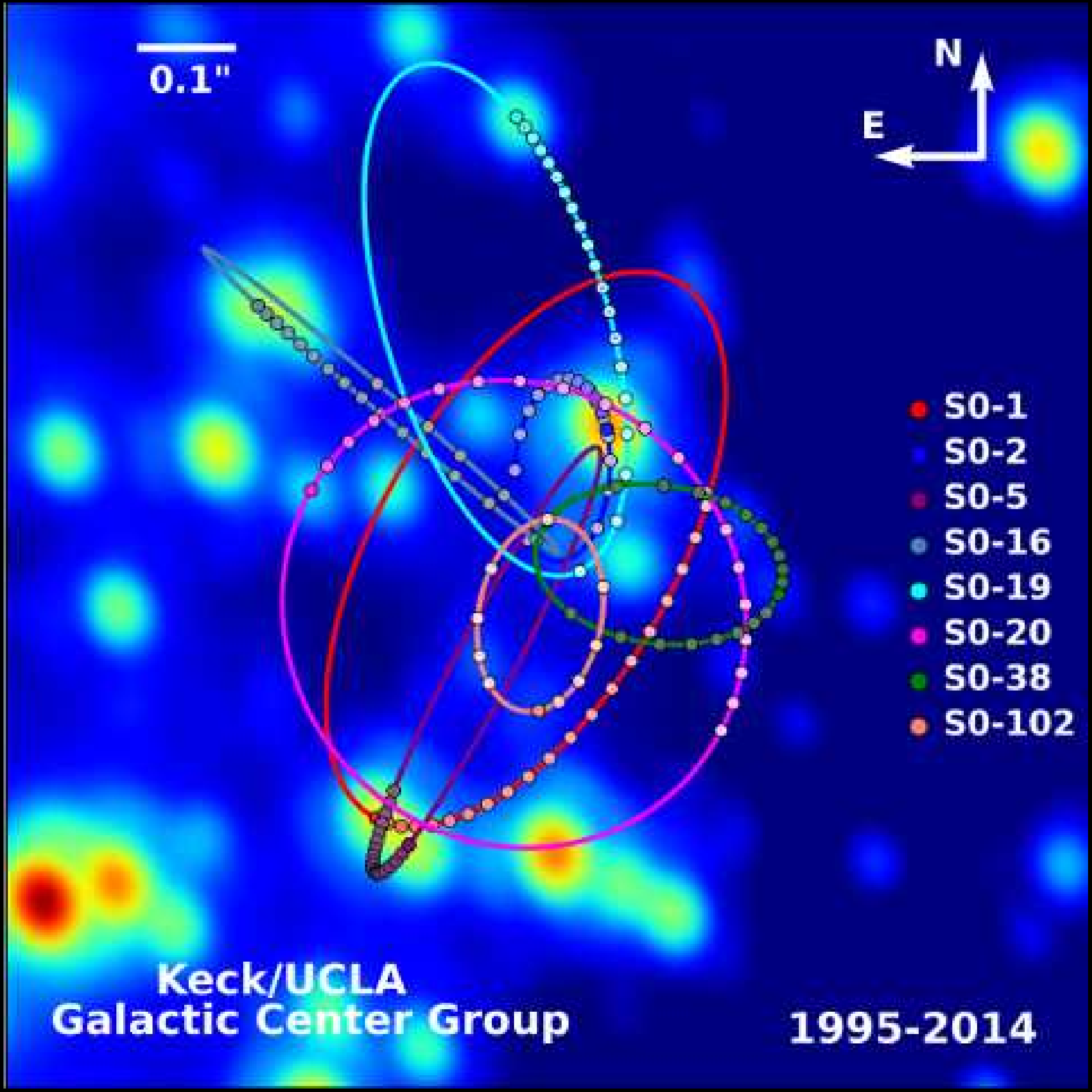


Chandra (X-ray), pink + VLA (radio), blue



Chandra (X-ray)

The Galactic Centre BH

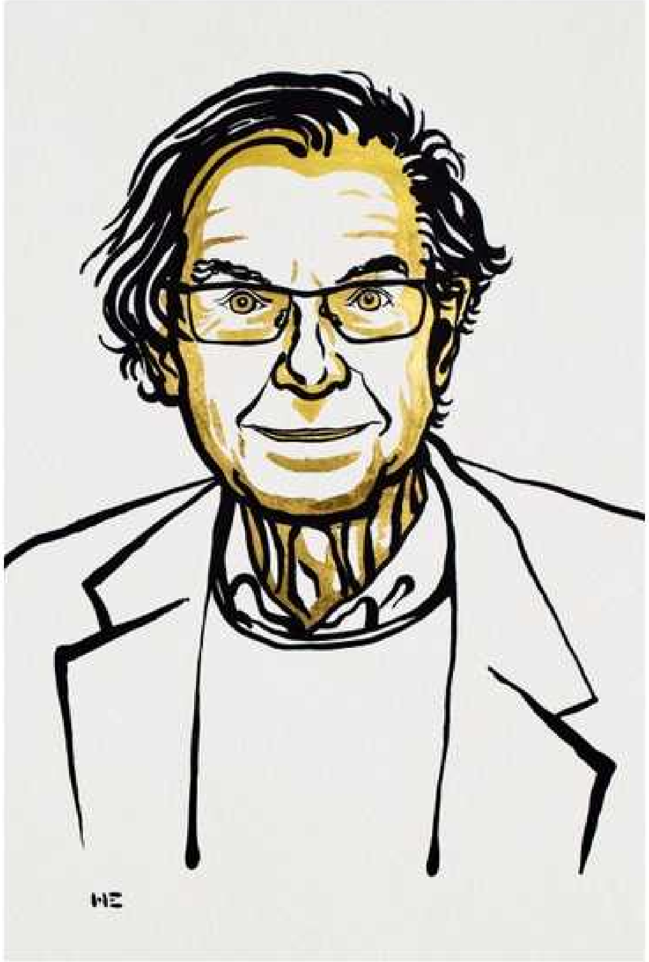


http://www.astro.ucla.edu/~ghezgroup/gc/blackhole.html

https://youtu.be/xHMZOaQttqw

https://youtu.be/if2opecmev8

The Nobel Prize in Physics 2020



Ill. Niklas Elmehed. © Nobel Media.

Roger Penrose

Prize share: 1/2

Ill. Niklas Elmehed. © Nobel Media.

Reinhard Genzel

Prize share: 1/4

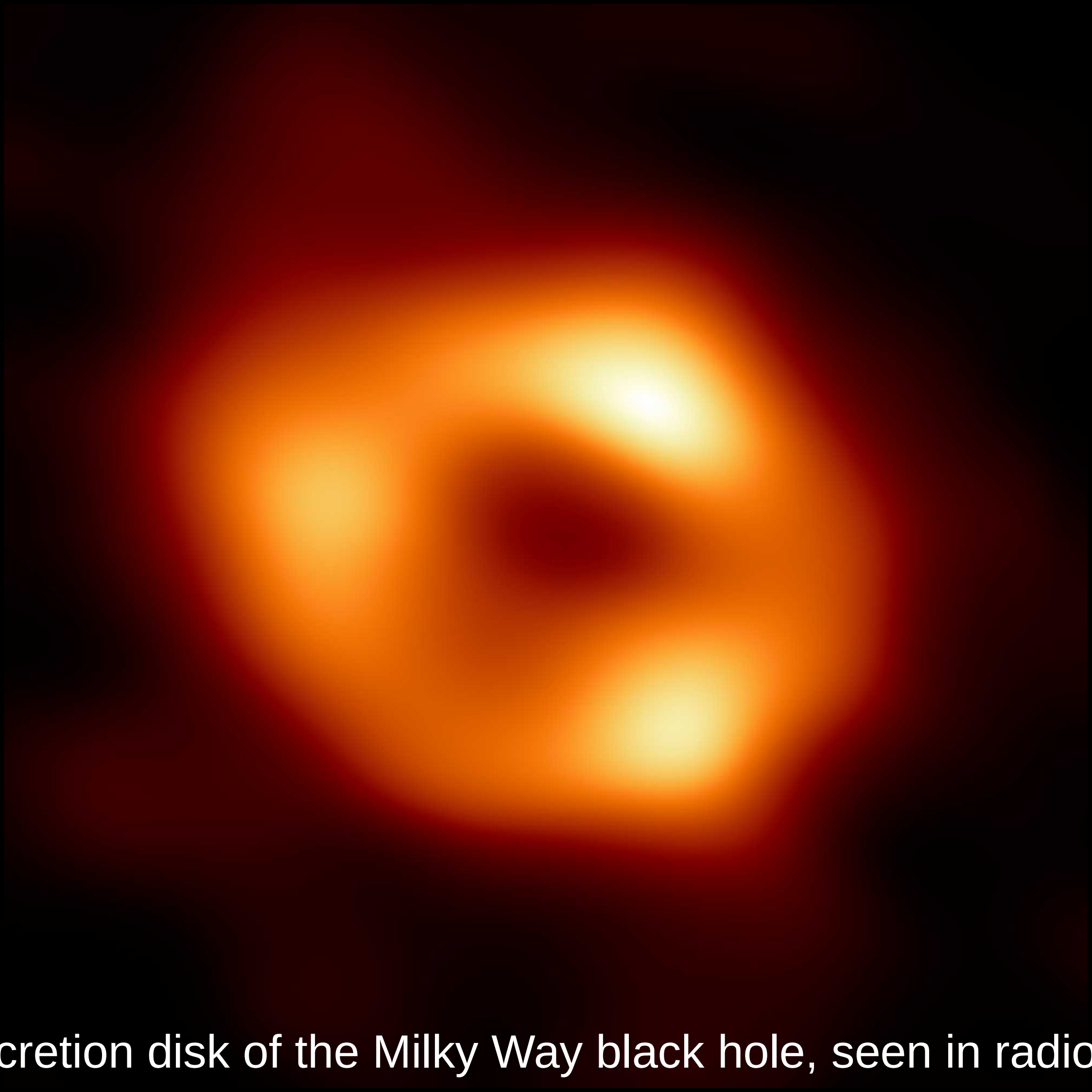
Ill. Niklas Elmehed. © Nobel Media.

Andrea Ghez

Prize share: 1/4

Event Horizon Telescope (EHT)

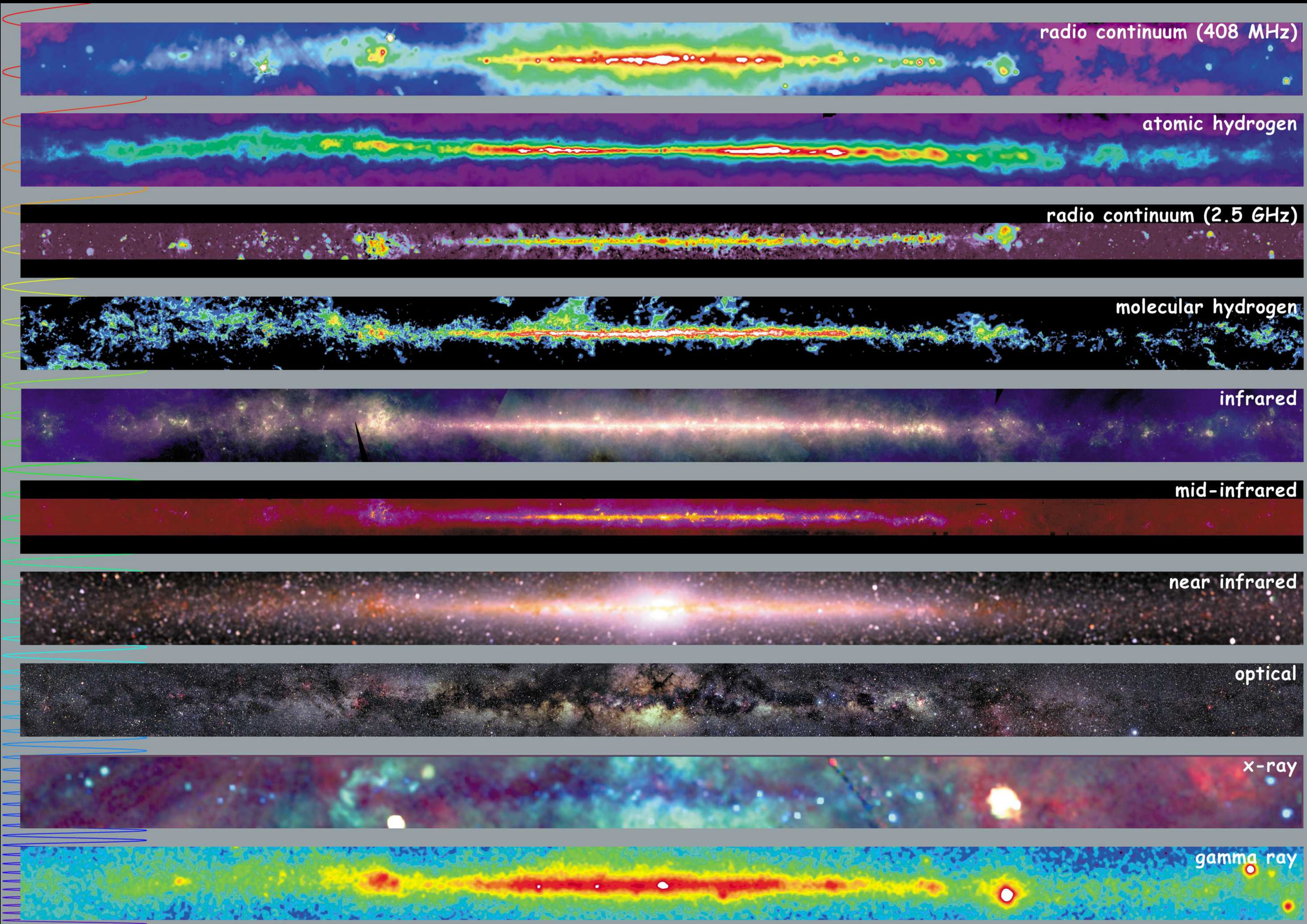
2019



The accretion disk of the Milky Way black hole, seen in radio

The Milky Way in different wavelength

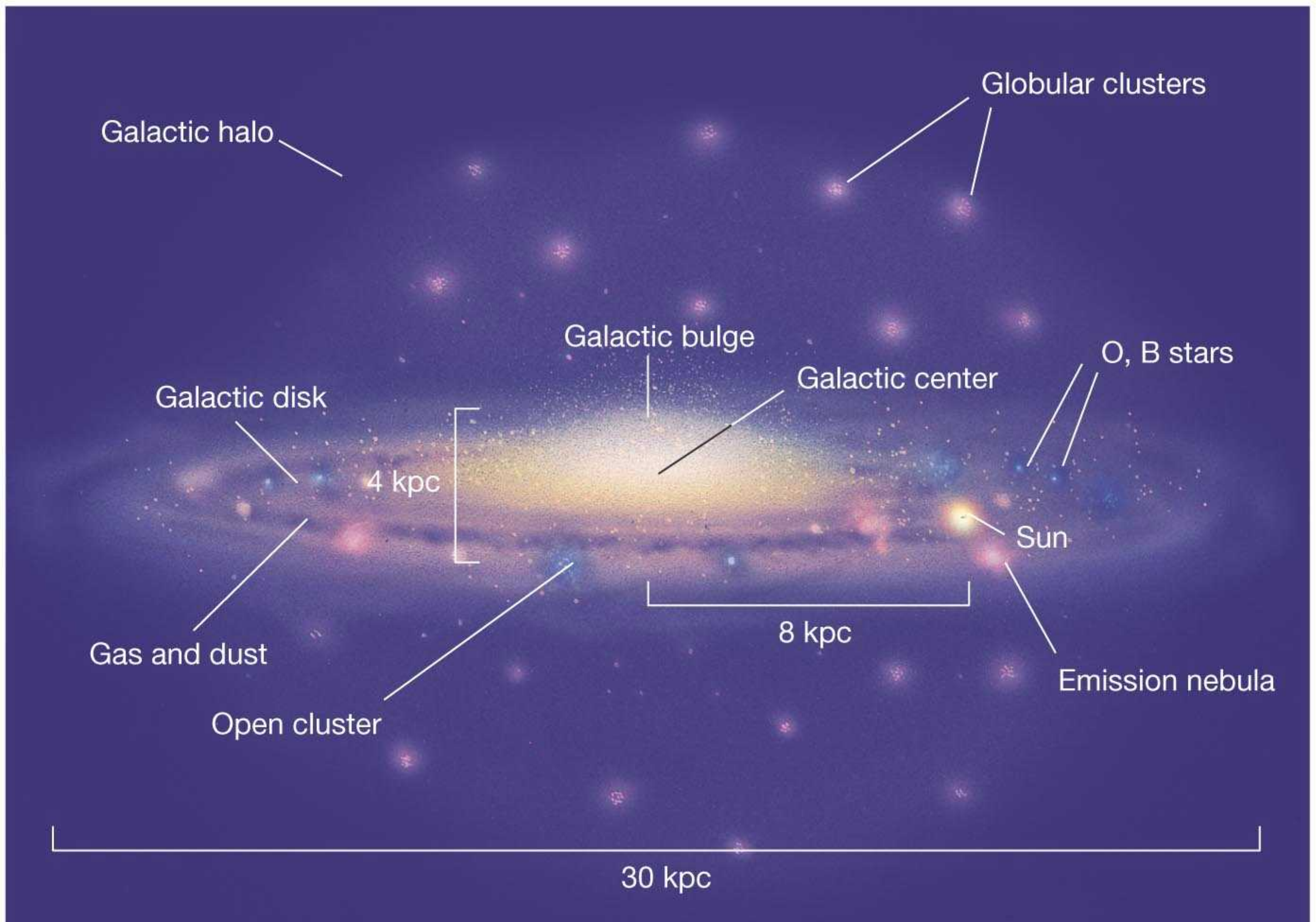
The Milky Way in different wavelength



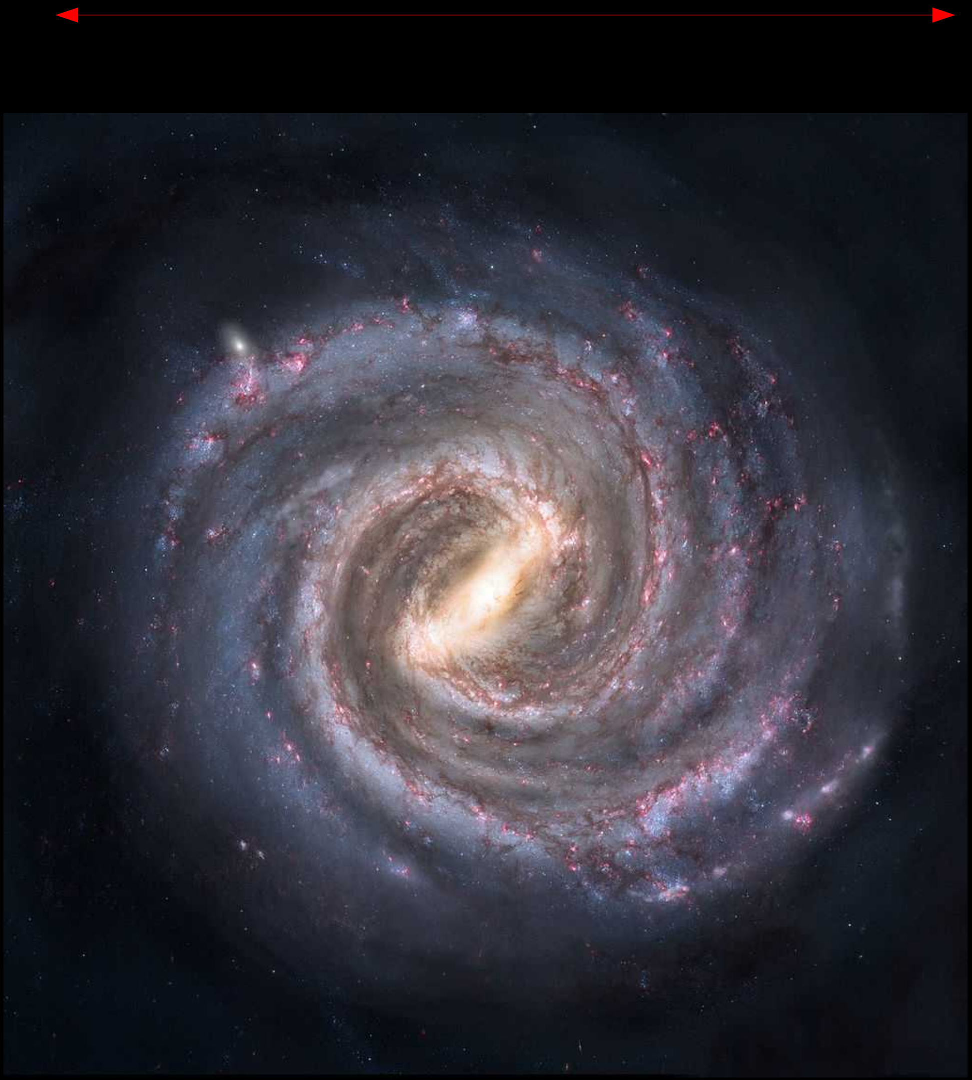
http://adc.gsfc.nasa.gov/mw

Multiwavelength Milky Way

Components of the WM



Components of the WM



Diameter :

30 kpc

Total mass:

10^{12} $M_\odot$

Rotation :

200 Myr (sun)
500 Myr (ext.)

Stellar component : bulge/bar

$0.5\times10^{10}\,M_\odot$

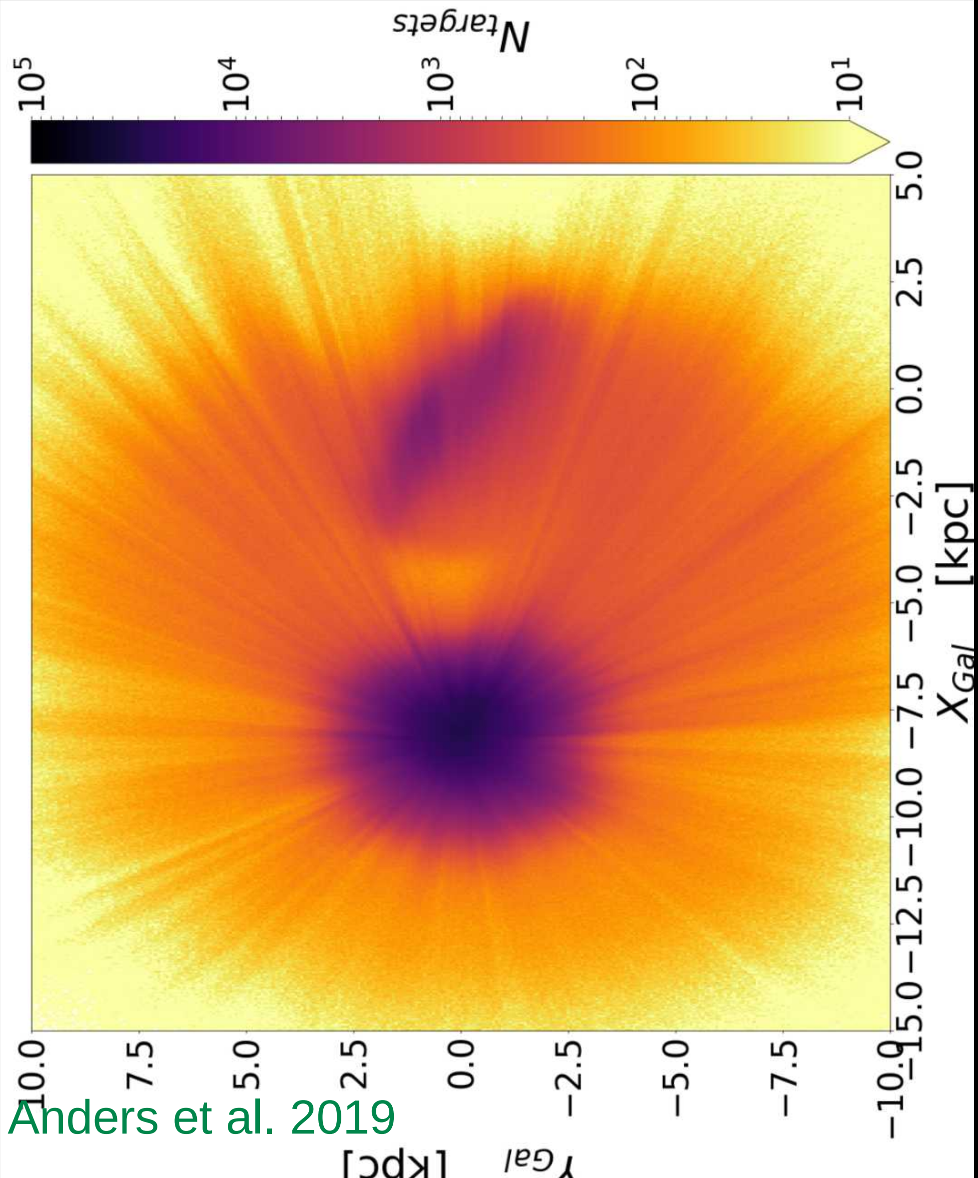
- old stars
- RMS vel ~150 km/s

Stellar component : bulge/bar

$0.5x10^{10} M_{\odot}$

265 millions of stars !

https://sci.esa.int/j/61461



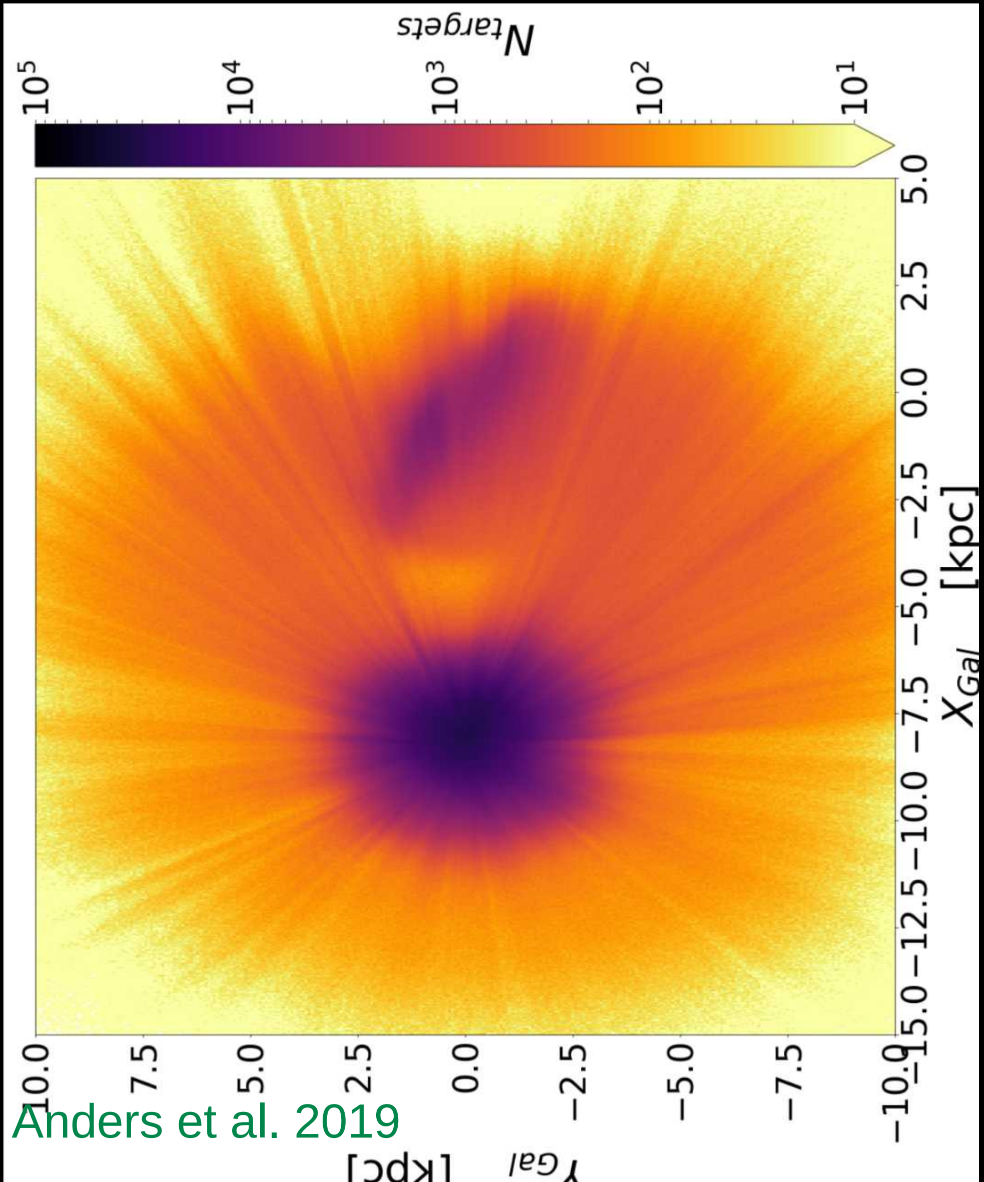
Anders et al. 2019

Stellar component : bulge/bar

$0.5\text{x}10^{10}\,M_\odot$

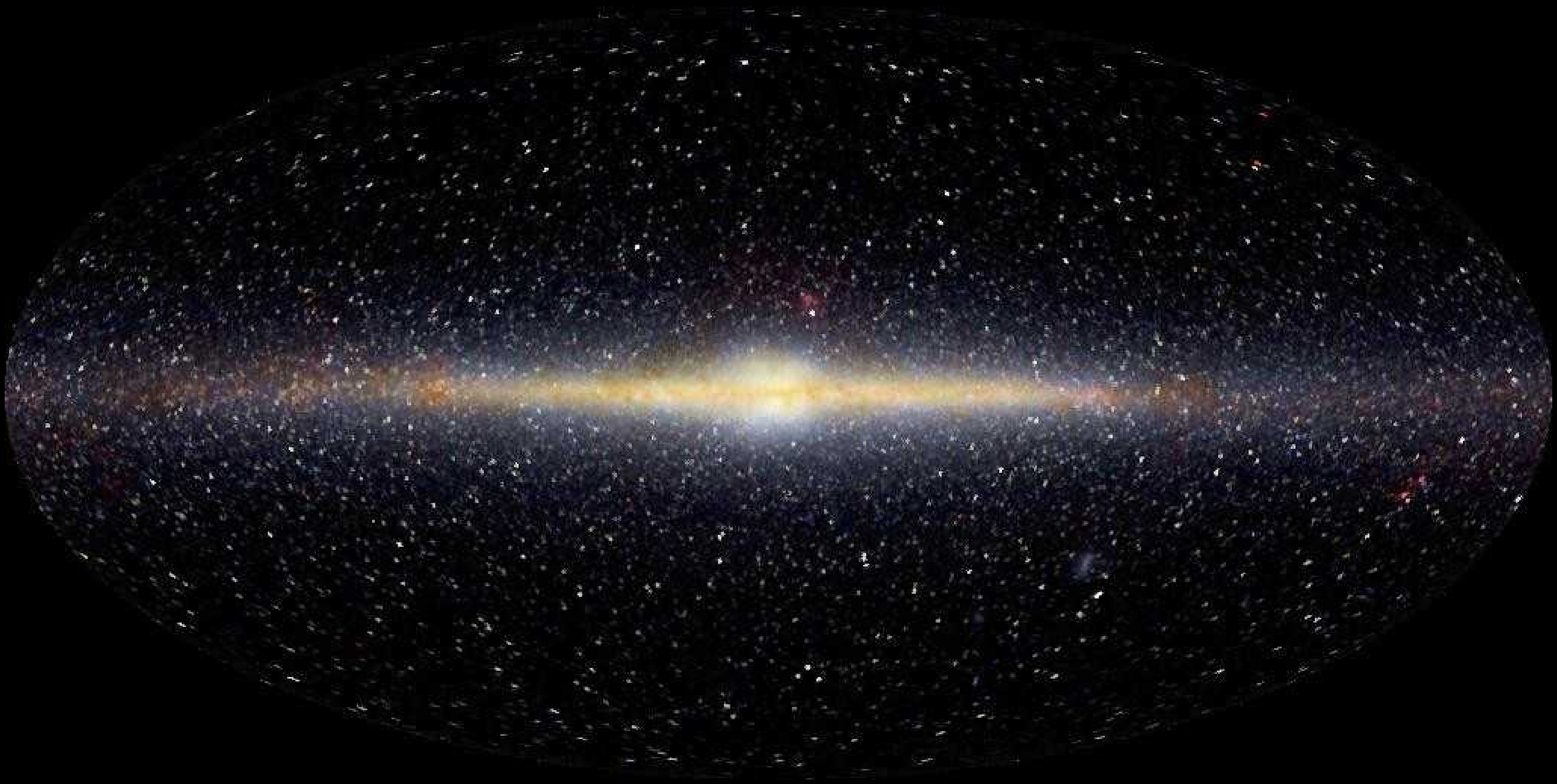
265 millions of stars !

https://sci.esa.int/j/61461



Anders et al. 2019

COBE satellite view of the MW in infrared light



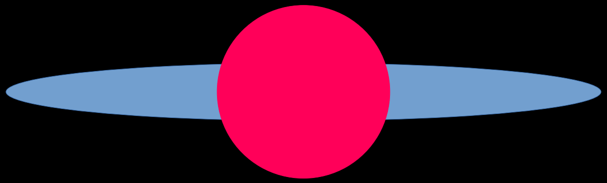
Robert Nemiroff (MTU) & Jerry Bonnell (USRA)

Stellar component : disk

$5 \times 10^{10} M_\odot$ (10 % of total)

thin disk:

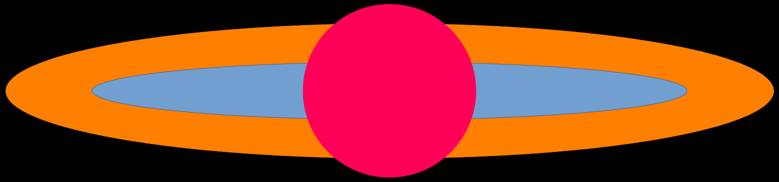
- 90% of the stellar disk
- scale height : ~300 pc
- RMS vel ~50 km/s



Stellar component : disk

$5x10^{10}$ $M_\odot$ (10 % of total)

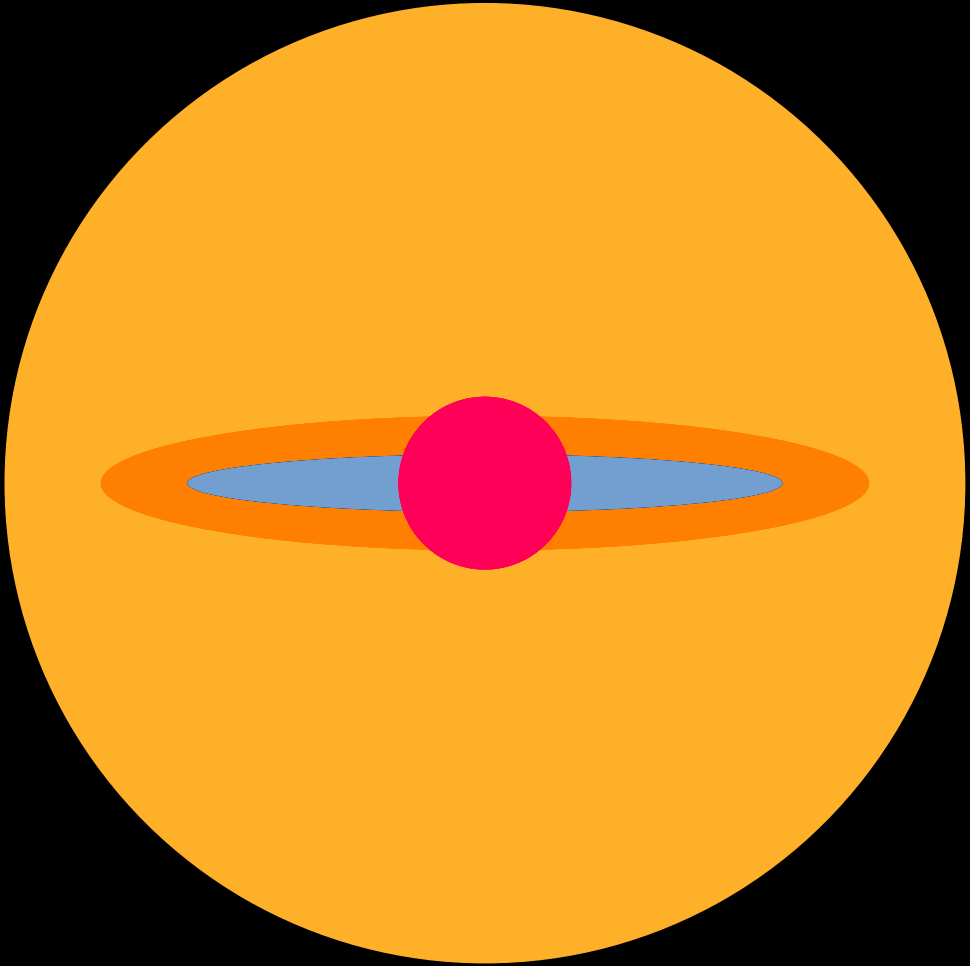
thick disk:
  - 10% of the stellar disk
  - scale height : ~1 kpc
  - RMS vel >~50 km/s



Stellar component : halo

$5 \times 10^8 \, M_\odot$ (1 % of stars)

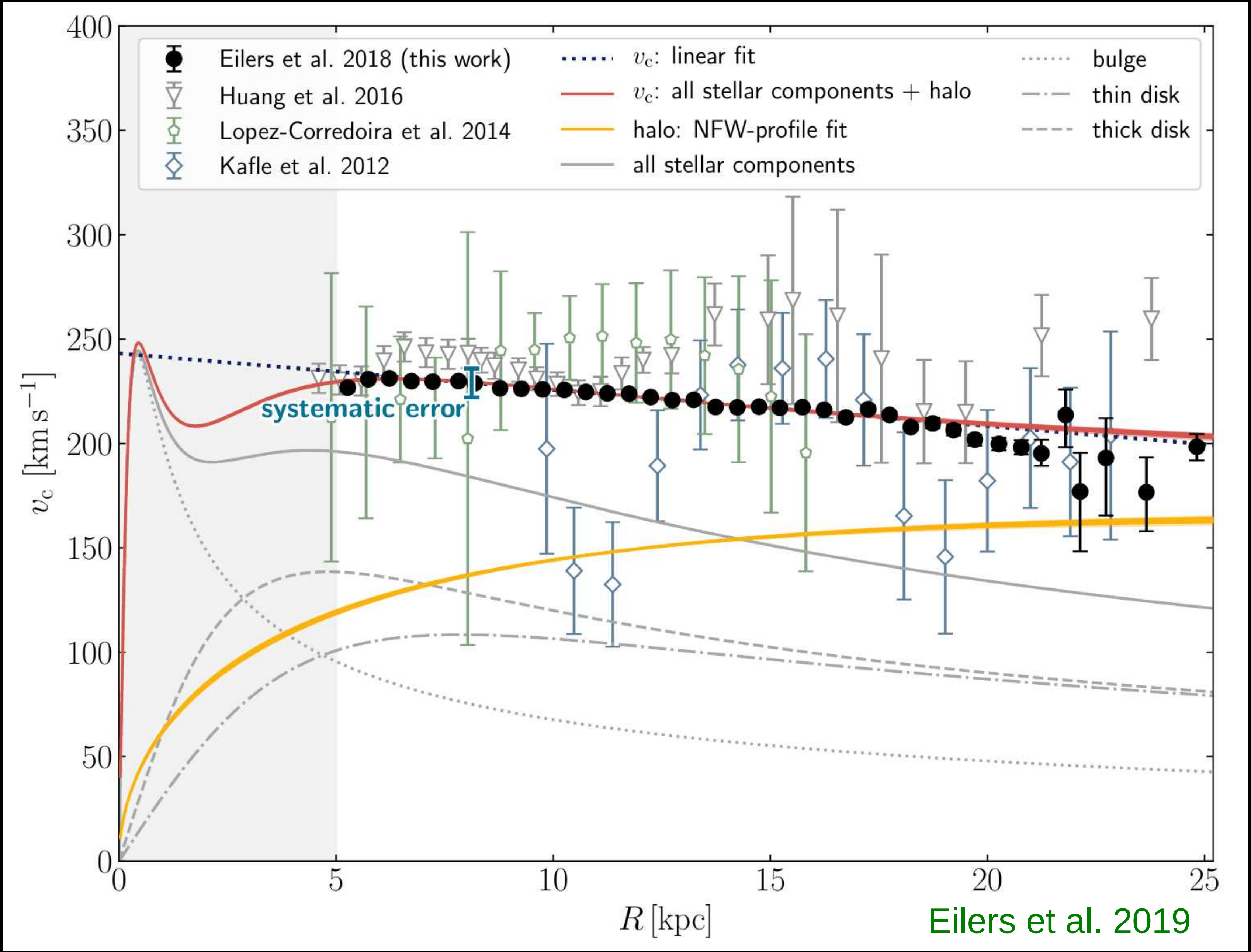
- old stars
- no mean rotation



Gaseous component : disk, HVC

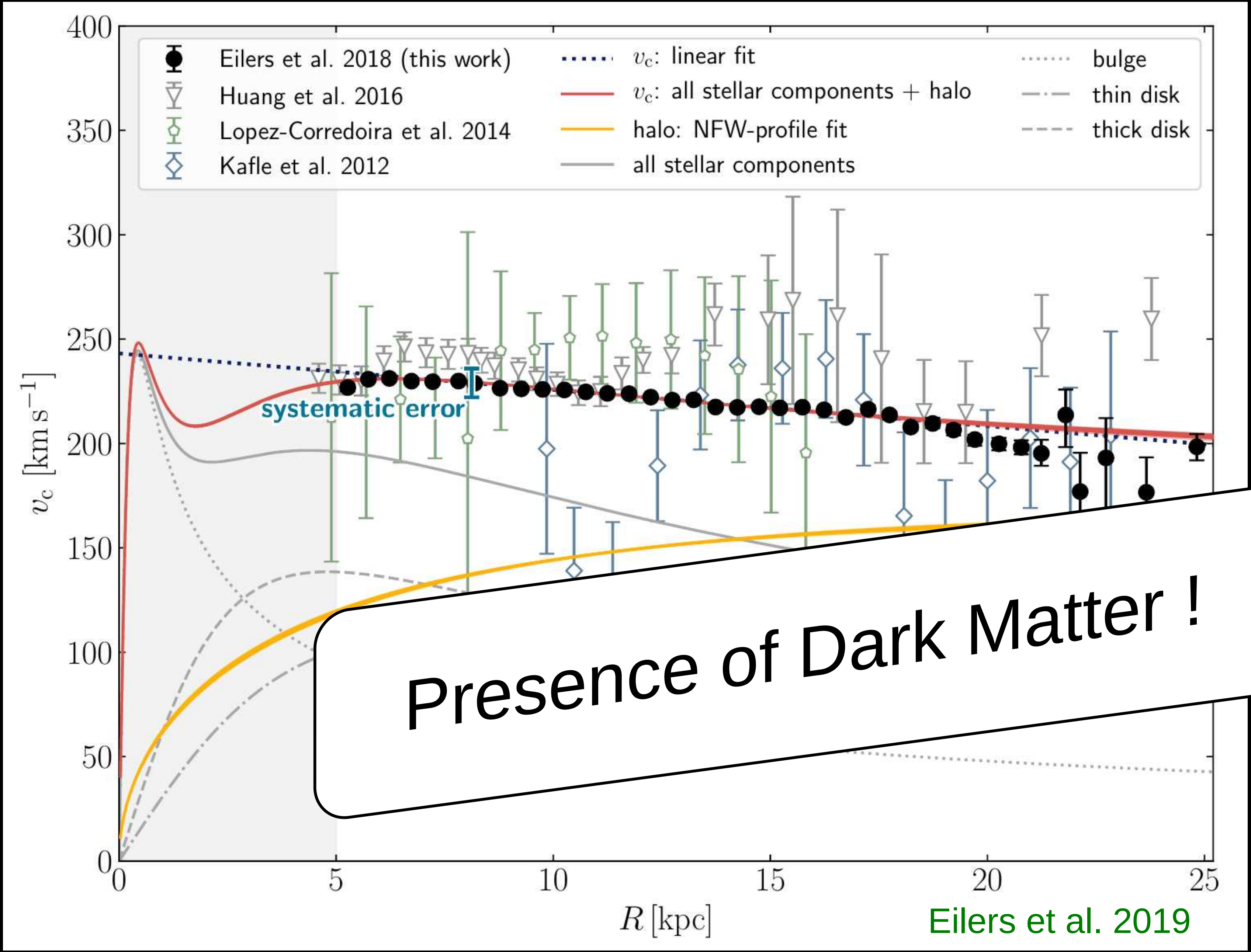
$10^9\ M_\odot$ (0.1 %)

The circular rotation curve of the MW



Eilers et al. 2019

The circular rotation curve of the MW

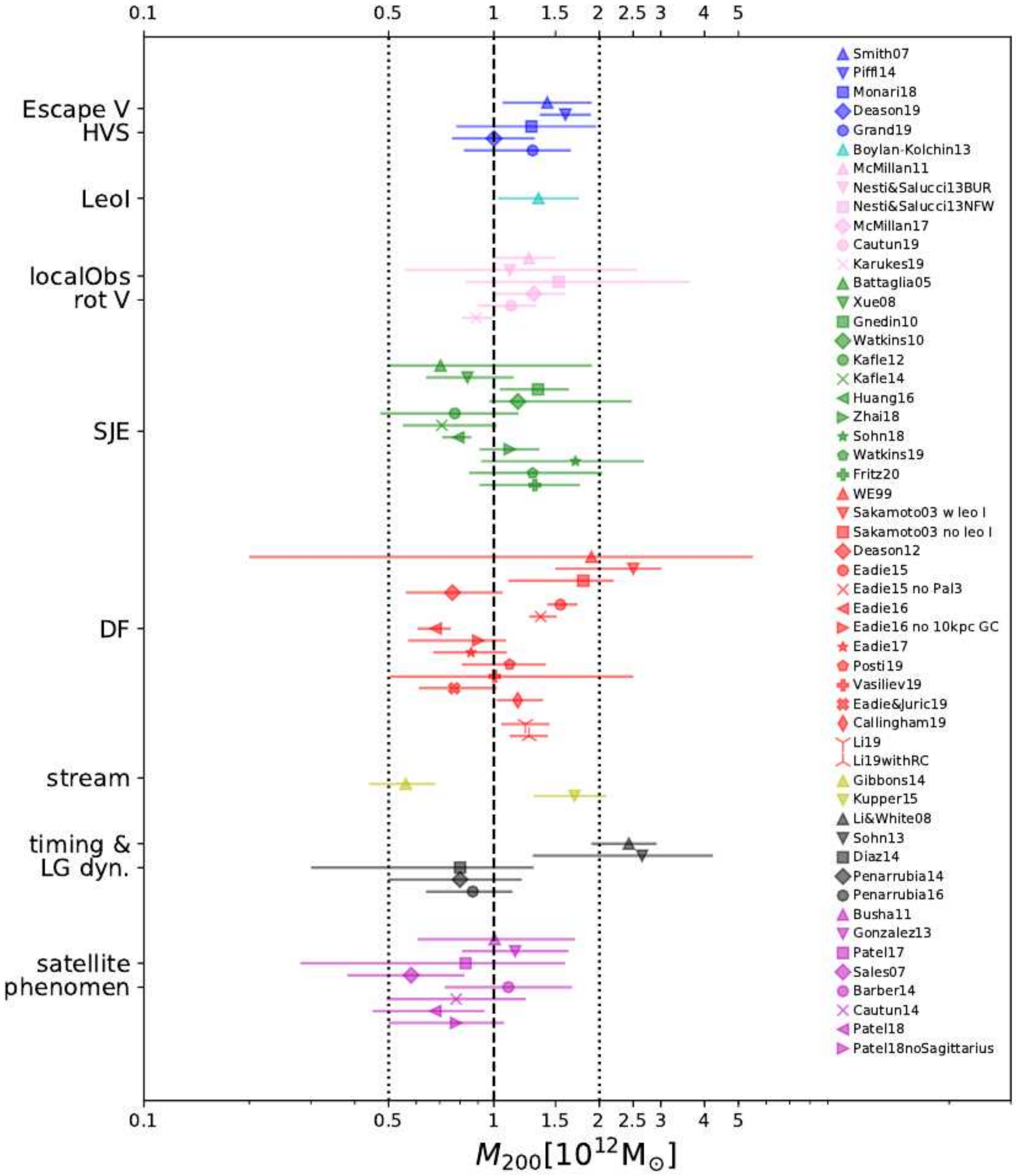


Presence of Dark Matter !

Eilers et al. 2019

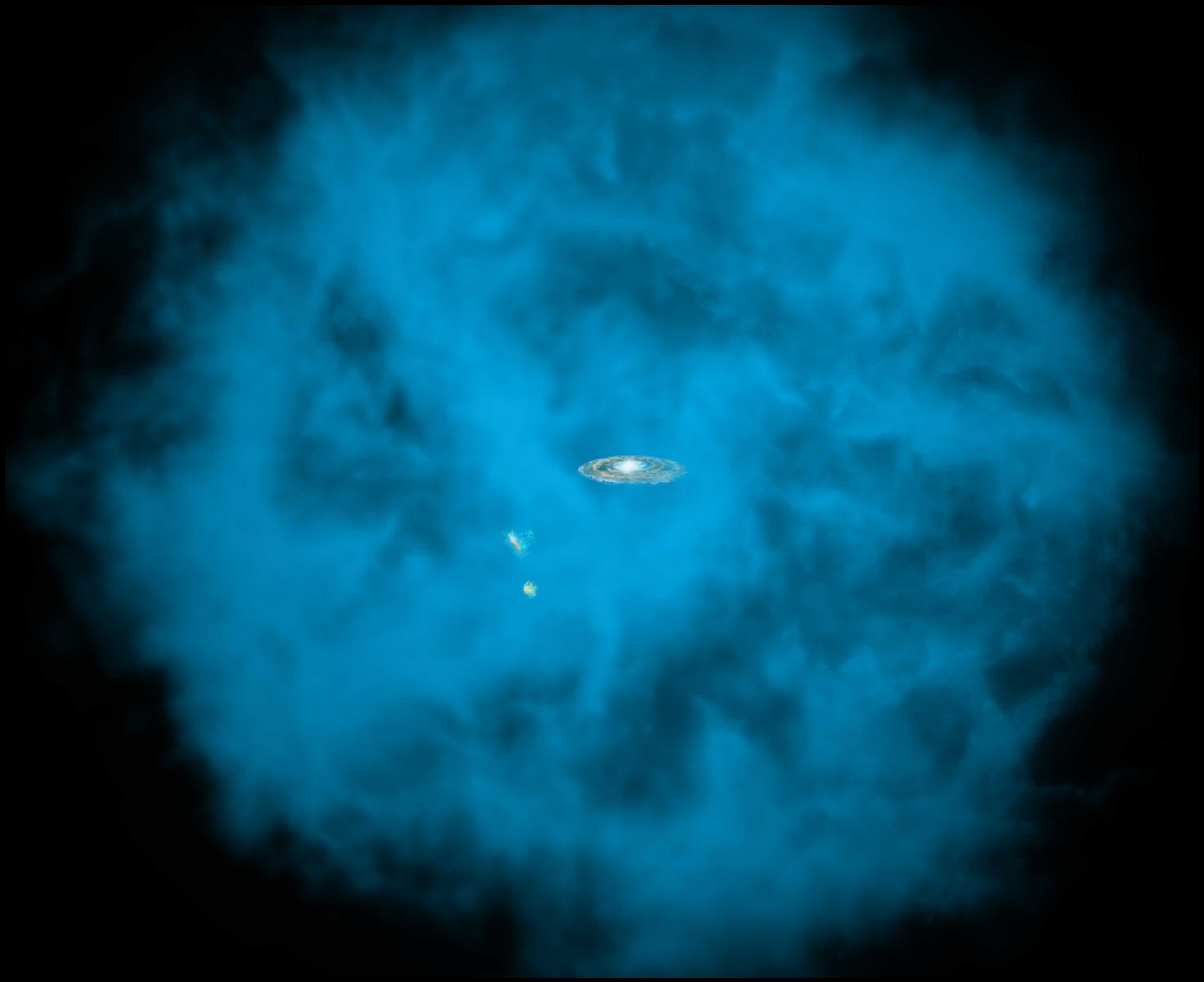
The Milky Way total (gravitational) mass

(Wang 2019, https://arxiv.org/abs/1912.02599)



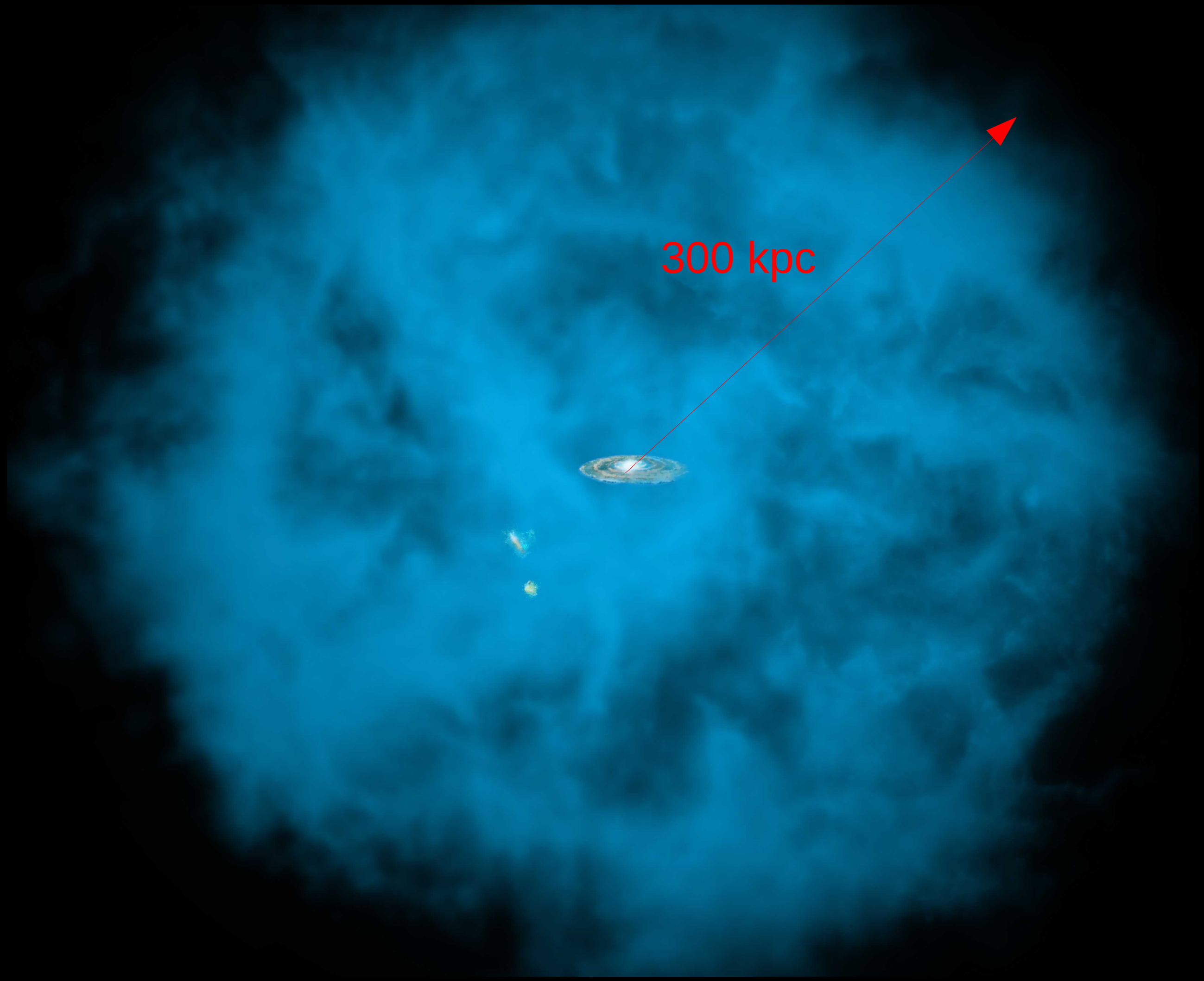
dark component : dark matter halo

about 90% of the total mass, $10^{12}\ M_\odot$



dark component : dark matter halo

about 90% of the total mass, 10^{12} $M_{\odot}$



The End